Astrophysics. Exam I Review

Chapter 1

- Periods: days and years.
- Physics I review: uniform circular motion, beat frequency
- Knowledge of the Ancient Greeks
- Stellar parallax
- Precession (what is it, period, consequences)
- Celestial Sphere
- Coordinates systems (altazimuth vs equatorial)
- Planetary configurations (opposition, etc)
- Synodic and sidereal periods; $1/S = 1/P_{in} - 1/P_{out}$

Chapter 2

- Ptolemaic system (epicycle, deferent, etc)
- Heliocentric system
- Copernicus, Tycho, Kepler, Galileo, Newton
- Ellipses ($r = \frac{a(\epsilon^2-1)}{1+\epsilon\cos\theta}$)
- Kepler's Laws
- Newton's fixes to Kepler's Laws
- Law of Universal Gravitation
- Coordinate conventions for 2-body problem . (Relative orbit or C.O.M,absolute coords)
- Reduced mass: $\mu = \frac{m_1 m_2}{m_1+m_2}$
- Center of mass (COM) coords: $\vec{r}_1 = -\frac{\mu}{m_1}\vec{r}$; and $\vec{r}_2 = \frac{\mu}{m_2}\vec{r}$

- Center of mass (COM) coords: $\vec{r} = \vec{r}_2 - \vec{r}_1$
- Total energy in terms of reduced mass:

$$E_{tot} = \frac{1}{2}\mu v^2 - G\frac{M\mu}{r}$$

- Total orbital angular momentum

$$\vec{L}_{tot} = \mu \vec{r} \times \vec{v}$$

- Results from the derivation of Kepler's 2nd law
  1. $\frac{d\vec{L}}{dt} = 0$ (angular momentum is constant in 2-body problem)
  2. $\frac{dA}{dt} = \frac{L}{2\mu}$
  3. $L = \mu\sqrt{GMa(1-e^2)}$
- The total energy of a 2-body system is 1/2 of the time-average potential energy: $E_{tot} = \frac{1}{2}\langle U \rangle$
- Escape velocity: $v_{esc} = \sqrt{2Gm/r}$
- Kepler's 3rd law (modified)

$$P^2 = \frac{4\pi^2}{G(m_1+m_2)}a^3$$

- Virial Theorem: for a multi-body system in equilibrium, the time-averaged kinetic energy and potential energy are related by:

$$-2\langle K \rangle = \langle U \rangle$$

- Also, for both multi-body systems and 2-body systems, total energy is:

$$\langle E \rangle = \frac{1}{2}\langle U \rangle$$

Chapter 3

- Parallax and distance. $d(pc) = \frac{1}{p''}$ (for baseline = 1 AU)
- Parallax (more general): $d = \frac{B}{2\tan p}$ (for baseline B)

- Proper motion ($\mu = v_\theta/d$ rad/s)
- Tangential velocity. v_θ(km/s)$= 4.74\mu d$ if μ in arcsec/yr, and d in pc
- Radial velocity ($v_r \sim c\frac{\Delta\lambda}{\lambda}$)
- True space motion ($v_T = \sqrt{v_r^2 + v_\theta^2}$)
- Flux, $F = \frac{L}{4\pi r^2}$ in Wm^{-2}
- Luminosity: total energy leaving an object in all directions over all wavelengths
- Monochromatic luminosity: $L_\lambda d\lambda$ = a luminosity only within the wavelength range λ to $\lambda + d\lambda$.
- Luminosity (blackbody) = $L = A\sigma T^4$.
- Luminosity (not quite perfect blackbody) = $L = \epsilon A\sigma T^4$.
- Magnitude System
  - 5 magnitudes difference correspondes to a flux ratio of 100X.
  - smaller numbers means brighter
  - apparent magnitude: $m = -2.5log_{10}\frac{F}{F_{ref}}$
  - absolute magnitude, M: the apparent magnitude of a star at the standard reference distance (10 pc = 32.6 ly).
  - absolute magnitude, $M = -2.5log_{10}\frac{L}{L_{ref}}$ (L_{ref} is about $80 \times L_\odot$.)
  - absolute magnitude is a measure of luminosity, apparent is a measure of brightness.
  - Examples: $M_\odot = 4.76$, $m_\odot = -26.7$, $L_\odot = 3.826 \times 10^{26}W$
  - Distance modulus, (m-M),: $(m - M) = 5log_{10}\frac{d}{10pc}$
  - Distance modulus: an alternative measure of distance that directly tells you how the brightness of the object differs from its brightness at 10 pc.
- Wave nature of light
  - Light has wave properties: interference pattern formed by double-slit
  - $c = \lambda\nu$
  - Poynting Vector: $\vec{S} = \frac{1}{\mu_0}\vec{E} \times \vec{B}$ a monochromatic flux
  - Time-averaged Poynting Vector: $\langle S\rangle = \frac{1}{2\mu_0}E_0B_0$ (mks)

- Radiation pressure is greater when light is completely reflected than when light is absorbed - transfer of momentum.
- Radiation pressure: $P_{rad} = \frac{S}{c}$
- Radiation force, pure absorption: $F_{rad} = \frac{SA}{c} \cos\theta$
- Radiation force, pure reflection: $F_{rad} = \frac{2SA}{c} \cos^2\theta$

- Blackbody radiation
  - Blackbody: an ideal emitter and absorber.
  - Blackbody absorption: 100%
  - Blackbody emission: spectrum obeys Planck function, Wien's Law, and the Stefan-Boltmann Law.
  - Wien's Law: $\lambda_{max}T = 0.0029mK = 2.9 \times 10^7 \mathring{A}K = (5000\mathring{A})(5800K)$
  - Stefan-Boltzmann law: $F_{surf} = \sigma T^4$
  - Specific intensity: $I_\lambda(x, y, z, \hat{r})$ = Energy/s/area/solid angle/wavelength at point (x,y,z) in direction $\hat{r}$.
  - Planck's Law: $B_\lambda(T) = \frac{2hc^2/\lambda^5}{e^{hc/\lambda kT}-1}$ is the I_λ emitted from a blackbody.
  - Planck's Law: $B_\nu(T) = \frac{2h\nu^3/c^2}{e^{h\nu/kT}-1}$ is the I_ν emitted from a blackbody.
  - $B_\lambda d\lambda = -B_\nu d\nu$ to convert

Chapter 4

- Special relativity: the physics of high speeds
- Lorentz factor: $\gamma = \frac{1}{\sqrt{1-v^2/c^2}}$
- $z = \Delta\lambda/\lambda$
- $z = \frac{v_r}{c}$ for low v_r
- $z = \sqrt{\frac{1+v_r/c}{1-v_r/c}} - 1$ for high v_r

Chapter 5 - POSTPONE TO EXAM 2