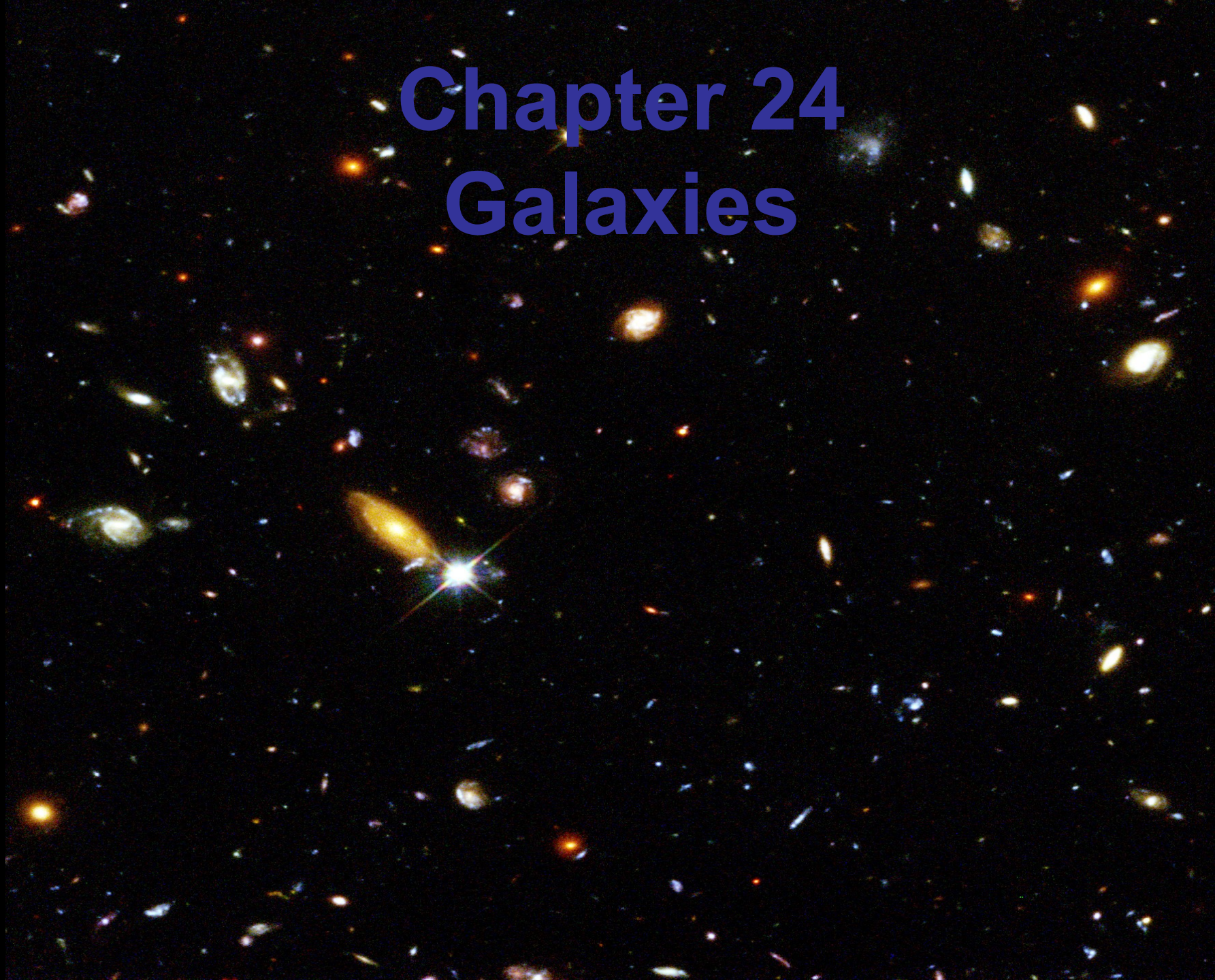


Chapter 24

Galaxies



Units of Chapter 24

24.1 Hubble's Galaxy Classification

24.2 The Distribution of Galaxies in Space

Tully-Fisher law, Cepheids

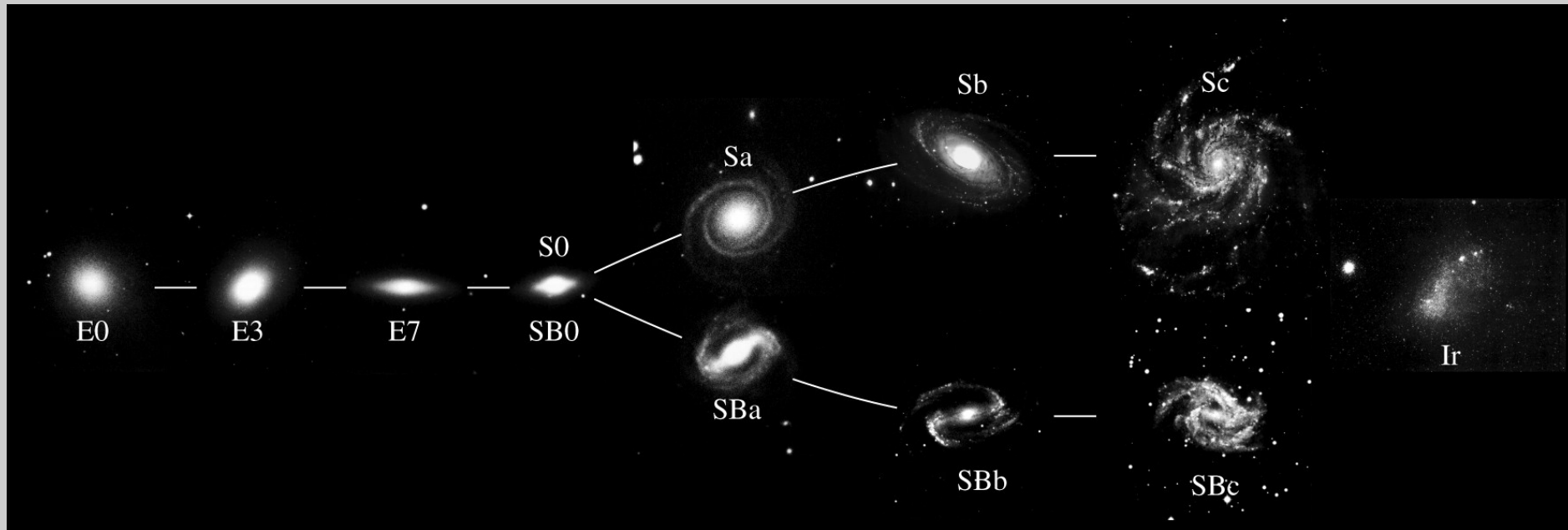
24.3 Hubble's Law

24.4 Active Galactic Nuclei (AGN)

24.5 Black Holes

24.1 Hubble's Galaxy Classification

Hubble “tuning-fork” diagram.



The Types:

E0, E1, ..., E7

S0, Sa, Sb, Sc

Irl, Irrl

SB0, Sba, Sbb, Sbc

“Early Types” <-----> “Late Types”

(Now we know evolution does not occur along this sequence.)

24.1 Hubble's Galaxy Classification

The Coma cluster of galaxies.



Clusters have many galaxies
at the same distance.

The wide range in galaxy sizes,
luminosities, colors, and
shapes is obvious.

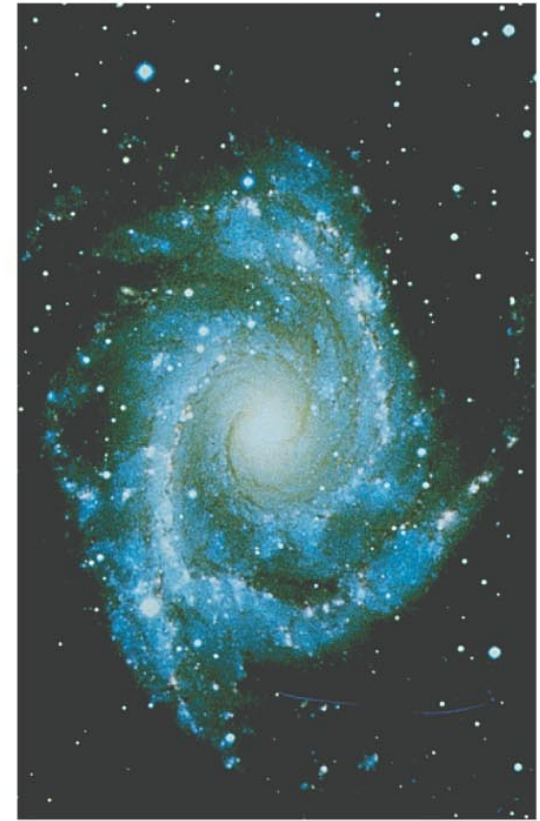
24.1 Hubble's Galaxy Classification



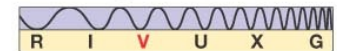
(a) M81 Type Sa



(b) M51 Type Sb



(c) NGC 2997 Type Sc



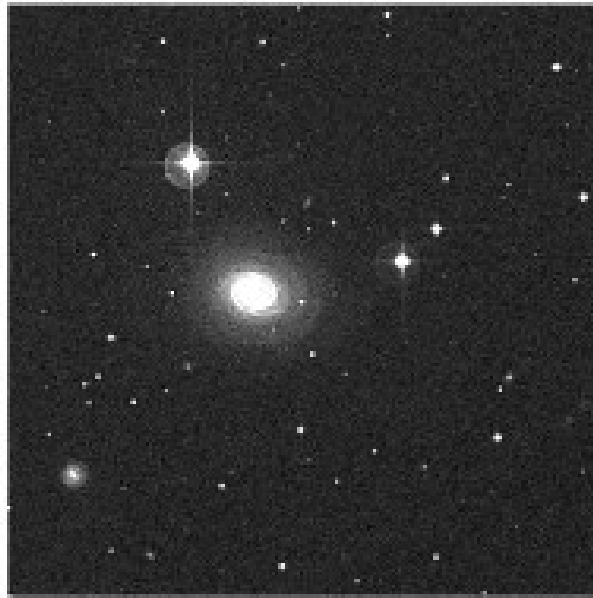
The components of spiral galaxies are the same as in our own galaxy: disk, halo, bulge, and spiral arms. Usually they also have nuclear clusters.

24.1 Hubble's Galaxy Classification

Spiral subtypes (a,b,c,d) are based on ...

- 1) the size of their central **bulge** relative to their **disk**. (Sa's have big bulges, Sc small.)
- 2) how tightly bound are the **spiral arms**. (Sa are tight, Sc loose).
- 3) how “clumpy” the arms appear. Types Sb and Sc are progressively clumpier.

Spirals



Sa

NGC 1357



Sb

M81



Sc

NGC 4321



a

M104



b

NGC 891



c

NGC 4631

Variety of Spiral Arms



a

M33

Flocculent spirals
(*fleecey*)



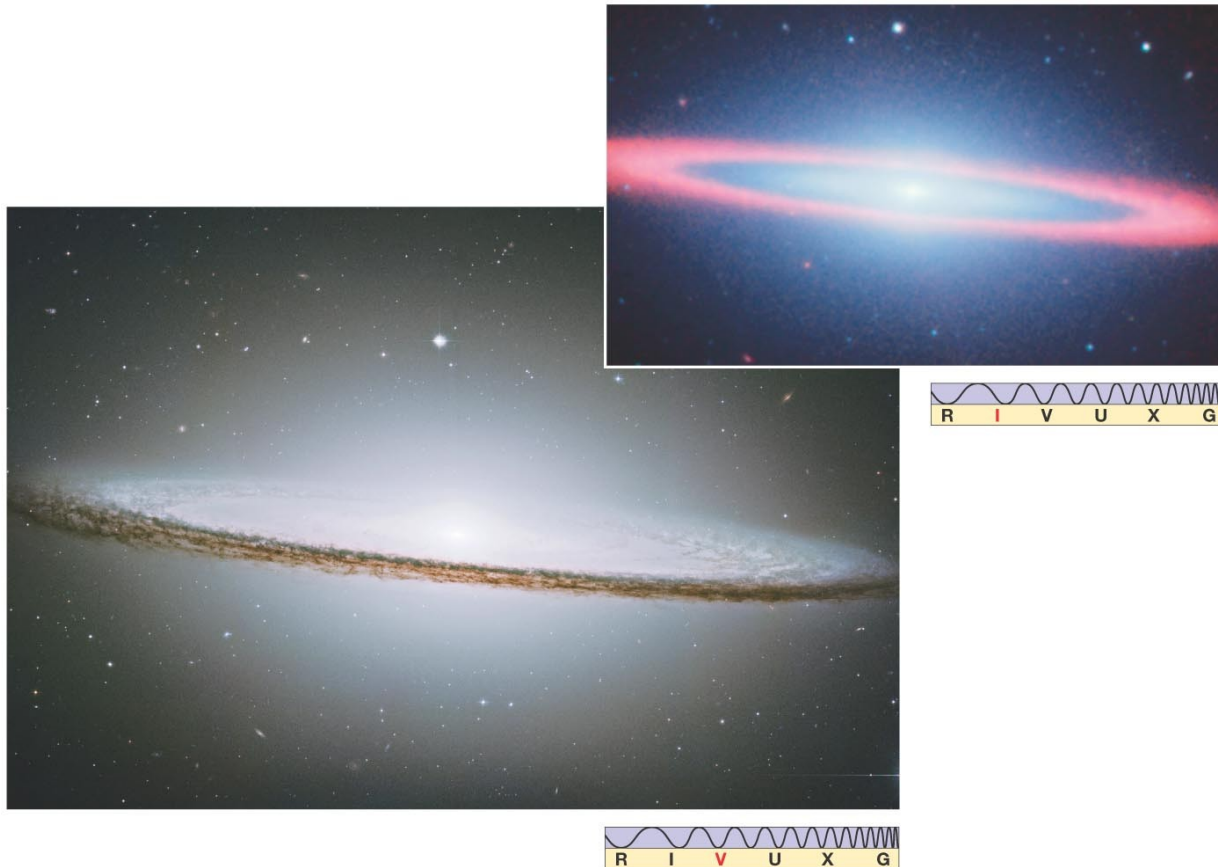
b

M74

Grand-design spirals
(*highly organized*)

24.1 Hubble's Galaxy Classification

The Sombrero galaxy, with its large central bulge, is a type Sa. Nearly edge-on disk is more well-defined by dust than stars.



24.1 Hubble's Galaxy Classification

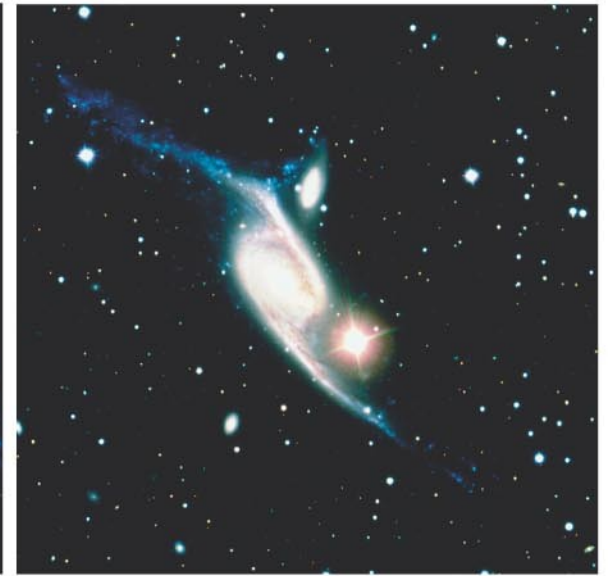
**Similar to the spiral galaxies are the
barred spirals:**



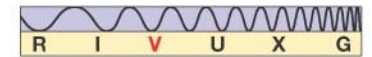
(a) NGC 1300 Type SBa



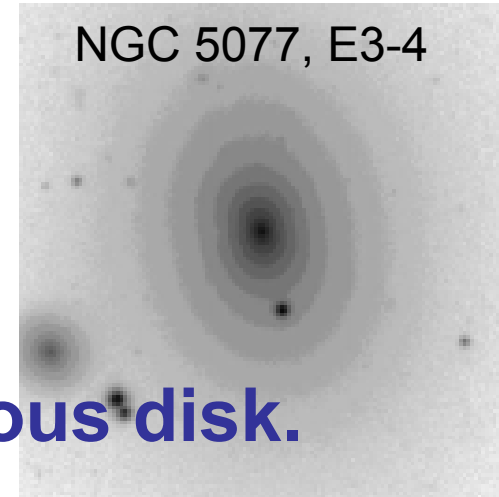
(b) NGC 1365 Type SBb



(c) NGC 6872 Type SBc



24.1 Hubble's Galaxy Classification



Elliptical galaxies

- **have no spiral arms and no obvious disk.**
- **come in many sizes, from giant ellipticals of trillions of stars, down to dwarf ellipticals of less than a million stars.**
- **contain very little cool gas and dust, and they show little evidence of ongoing star formation.**
- **Many are immersed in a large cloud of hot gas, visible in x-rays**

24.1 Hubble's Galaxy Classification

Elliptical subtypes are classified according to their shape from E0 (almost spherical) to E7 (the most elongated):



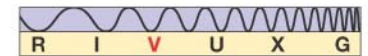
(a) M49 Type E2



(b) M84 Type E3

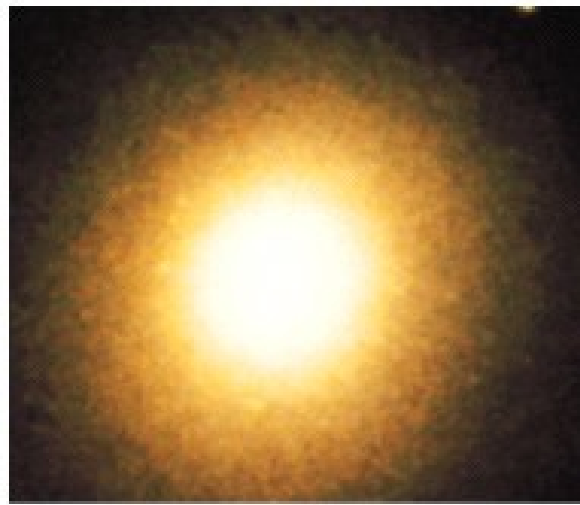


(c) M110 Type E5



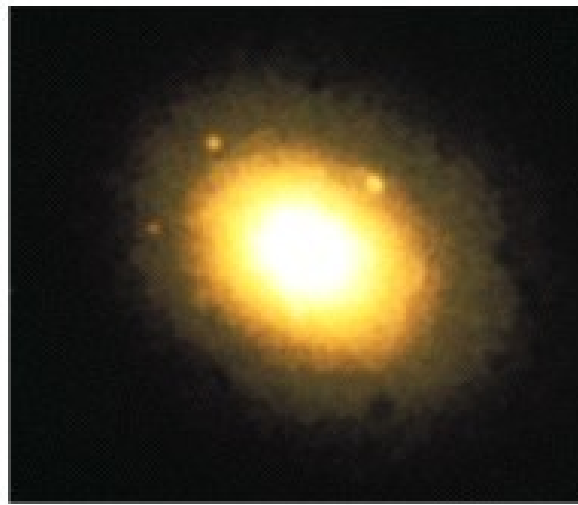
$$\text{Numerical class} = 10(\text{ellipticity}) = 10 (1 - b/a)$$

Elliptical galaxies



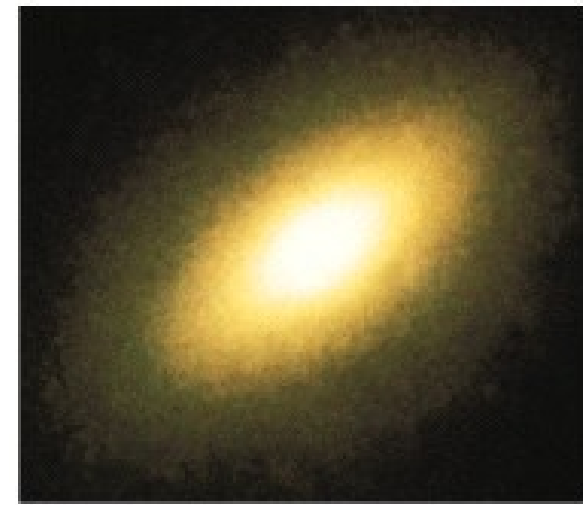
E0

M105



E3

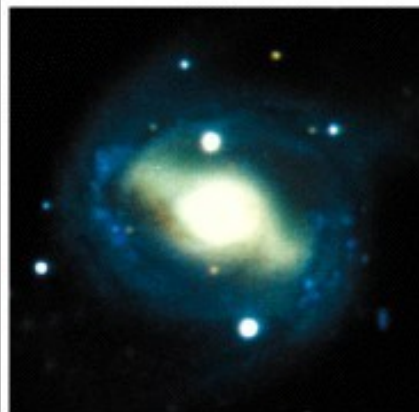
NGC 4365



E6

NGC 3377

Barred Spirals



SBa

NGC 4650



SBb

M83



SBc

NGC 1365

24.1 Hubble's Galaxy Classification

S0 (lenticular) and **SB0** galaxies have a disk and bulge, but no spiral arms and little interstellar gas:



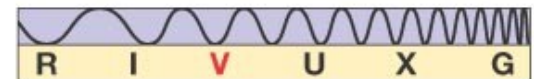
(a) NGC 1201

Type S0



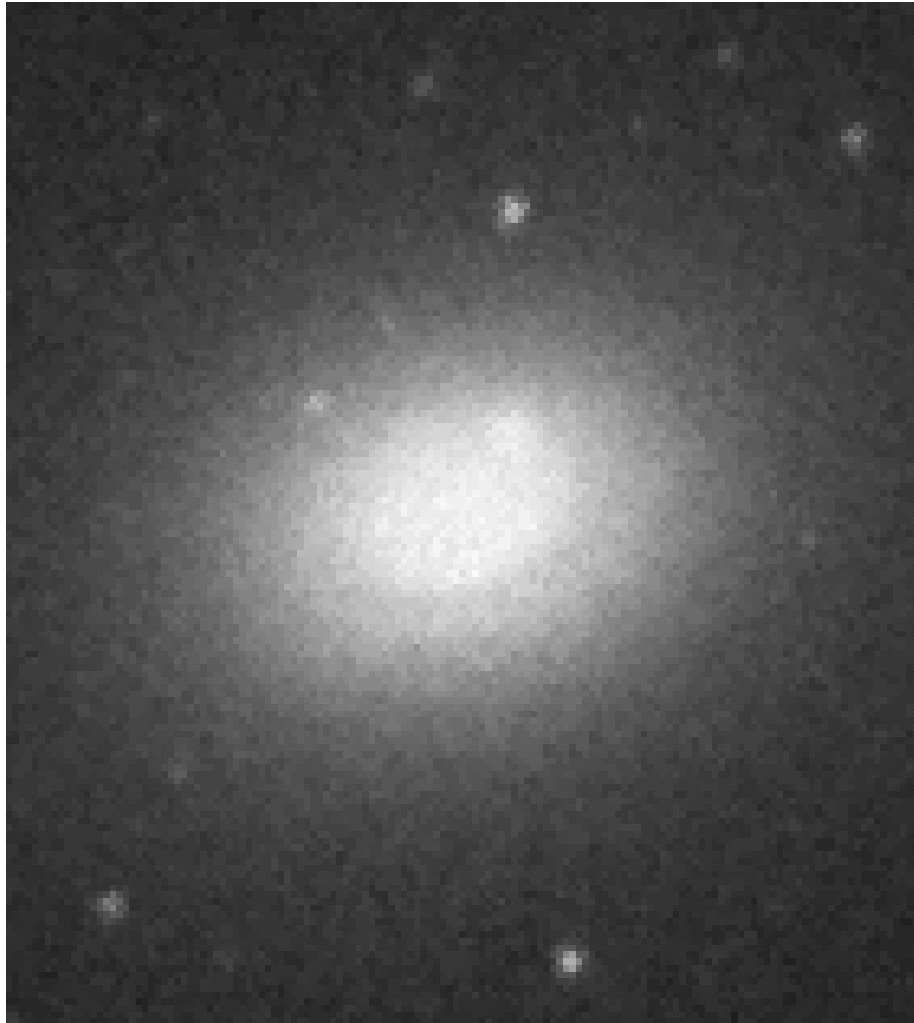
(b) NGC 2859

Type SB0



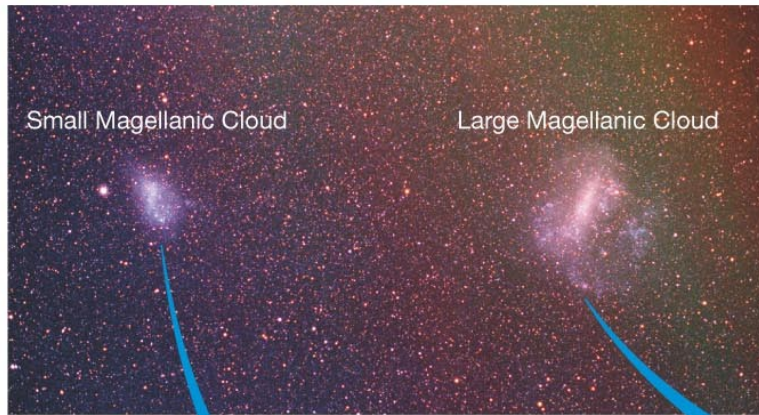
24.1 Hubble's Galaxy Classification

NGC 7049 (S0) ground-based vs Hubble



24.1 Hubble's Galaxy Classification

The irregular galaxies have a wide variety of shapes. The small and large Magellanic Clouds are prototypes for Type I (or Magellanic type) irregulars.



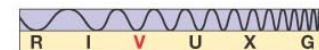
(a)



(b)



(c)

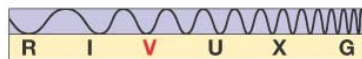


24.1 Hubble's Galaxy Classification

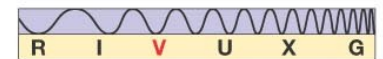
Here are two contrasting **irregular** galaxies:
Type I on the left, and Type II (smooth) on the right.



(a) NGC 1427A



(b) M82



24.1 Hubble's Galaxy Classification

The companion to
M51 is an Irr 2.



24.1 Hubble's Galaxy Classification

A summary of galaxy properties by type:

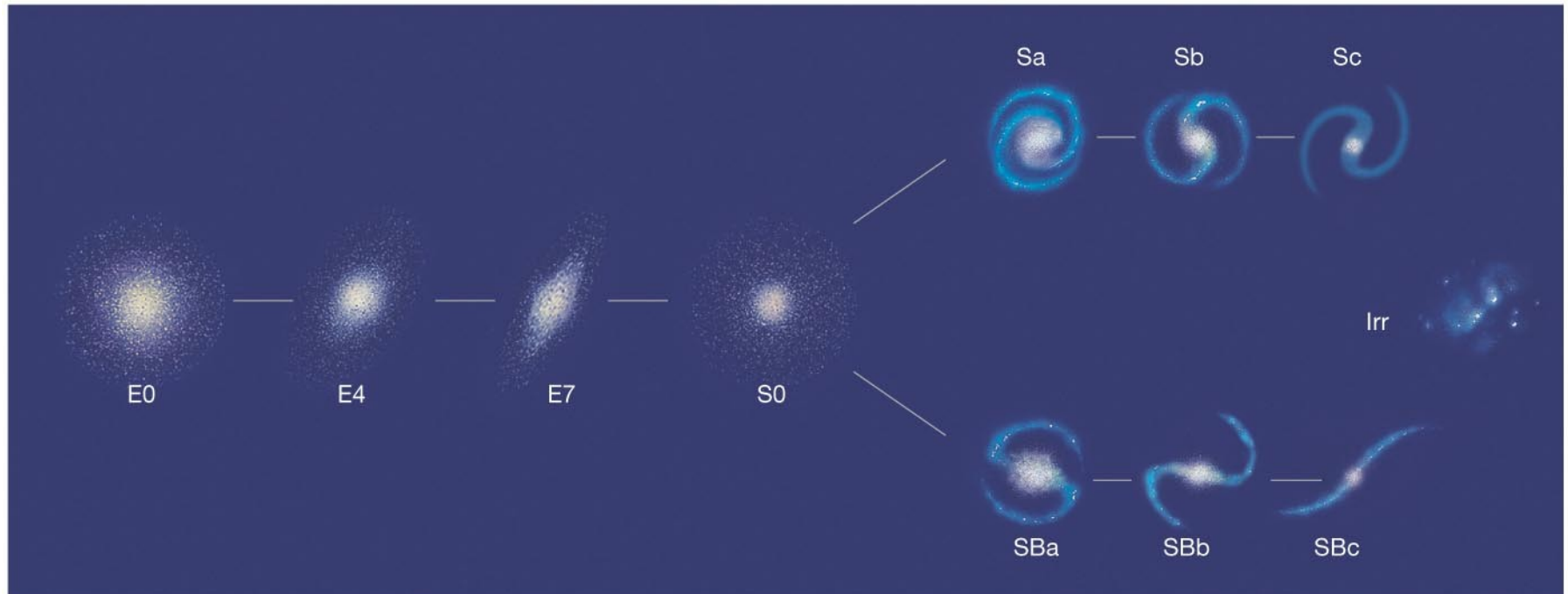
TABLE 24.1 Galaxy Properties by Type

	Spiral/Barred Spiral (S/SB)	Elliptical* (E)	Irregular (Irr)
Shape and structural properties	Highly flattened disk of stars and gas, containing spiral arms and thickening central bulge. Sa and SBa galaxies have the largest bulges, the least obvious spiral structure, and roughly spherical stellar halos. SB galaxies have an elongated central “bar” of stars and gas.	No disk. Stars smoothly distributed through an ellipsoidal volume ranging from nearly spherical (E0) to very flattened (E7) in shape. No obvious substructure other than a dense central nucleus.	No obvious structure. Irr II galaxies often have “explosive” appearances.
Stellar content	Disks contain both young and old stars; halos consist of old stars only.	Contain old stars only.	Contain both young and old stars.
Gas and dust	Disks contain substantial amounts of gas and dust; halos contain little of either.	Contain hot X-ray emitting gas, little or no cool gas and dust.	Very abundant in gas and dust.
Star formation	Ongoing star formation in spiral arms.	No significant star formation during the last 10 billion years.	Vigorous ongoing star formation.
Stellar motion	Gas and stars in disk move in circular orbits around the galactic center; halo stars have random orbits in three dimensions.	Stars have random orbits in three dimensions	Stars and gas have highly irregular orbits.

** As noted in the text, some giant ellipticals appear to be the result of collisions between gas-rich galaxies and are exceptions to many of the statements listed here.*

24.1 Hubble's Galaxy Classification

Hubble's “**tuning fork**” is a logical ordering of the galaxy classifications. But galaxies do **NOT** evolve along the fork.



See my “Galaxies along the Hubble Sequence” web page.

Galaxy Evolution

Classification in science is important because it can tell us about the nature of things and how they EVOLVE. Astronomers know:

- **Galaxies do not really evolve along the Hubble Sequence, although spirals can combine to form Irr 2's, peculiars, and eventually E's.**
- **More irregular, clumpy galaxies are seen at high redshifts (back in time).**
- **Red ellipticals exist way back to ~10 Gyrs ago.**
- **More ellipticals are found in denser regions.**
- **Quasars at $z=2-3$ and black holes play a role.**

Galaxy Evolution

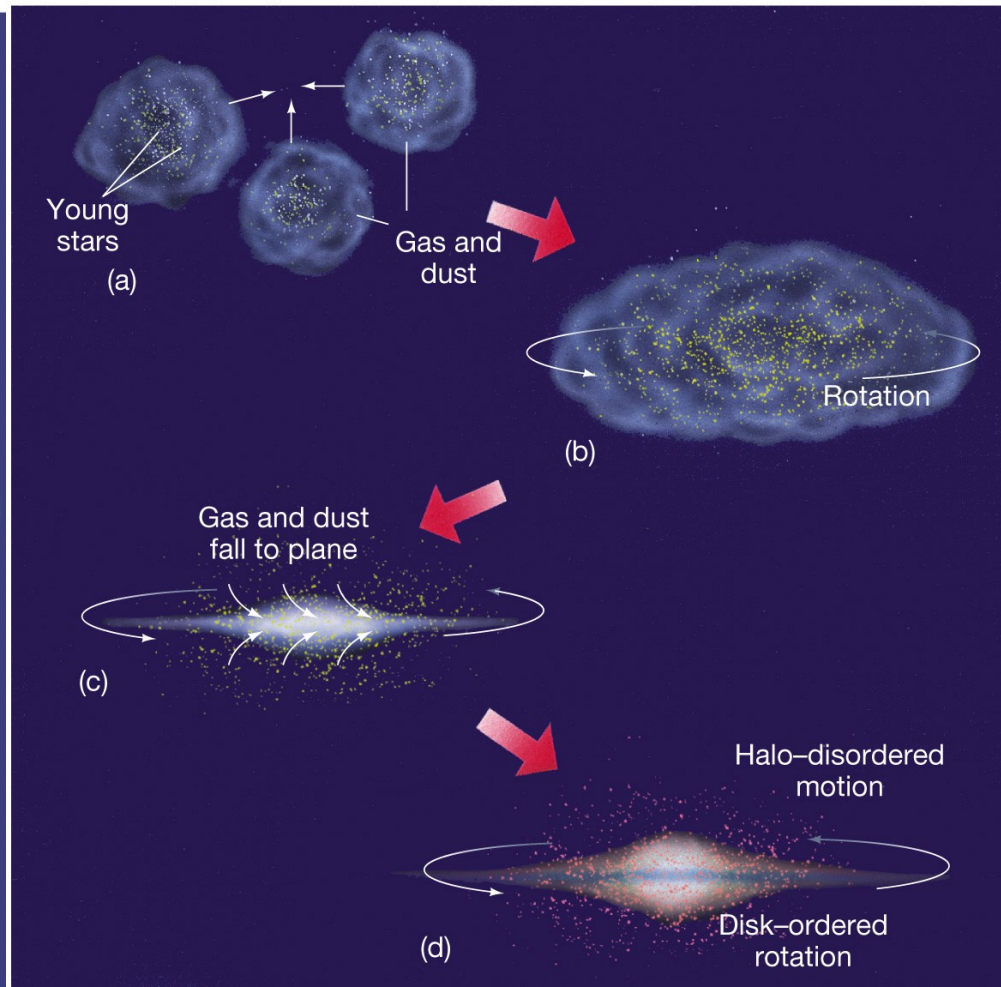
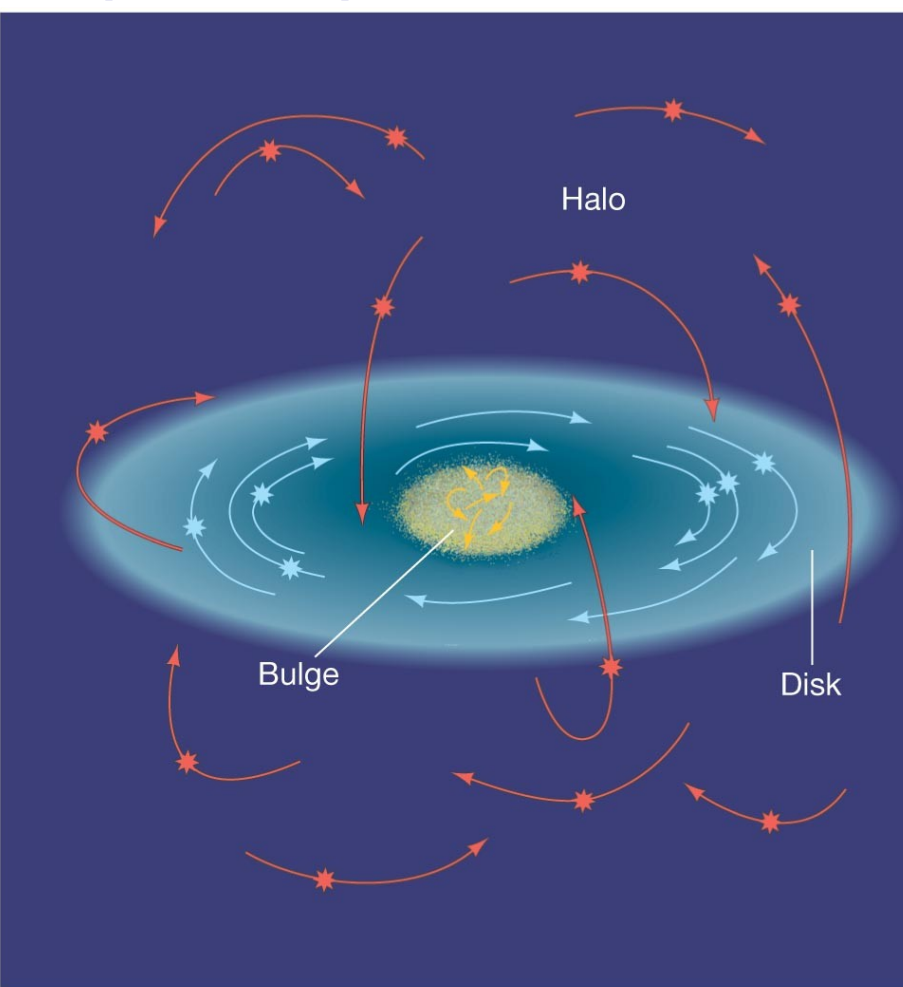
How did our galaxy (the Milky Way) form?

The following questions point to a theory...

- Where are the oldest (Pop II) and youngest (Pop I) stars found?
- How do the kinematics of the youngest and oldest stars differ?
- What can you say about the order of formation of the halo stars (and globular clusters), the disk, the bulge?

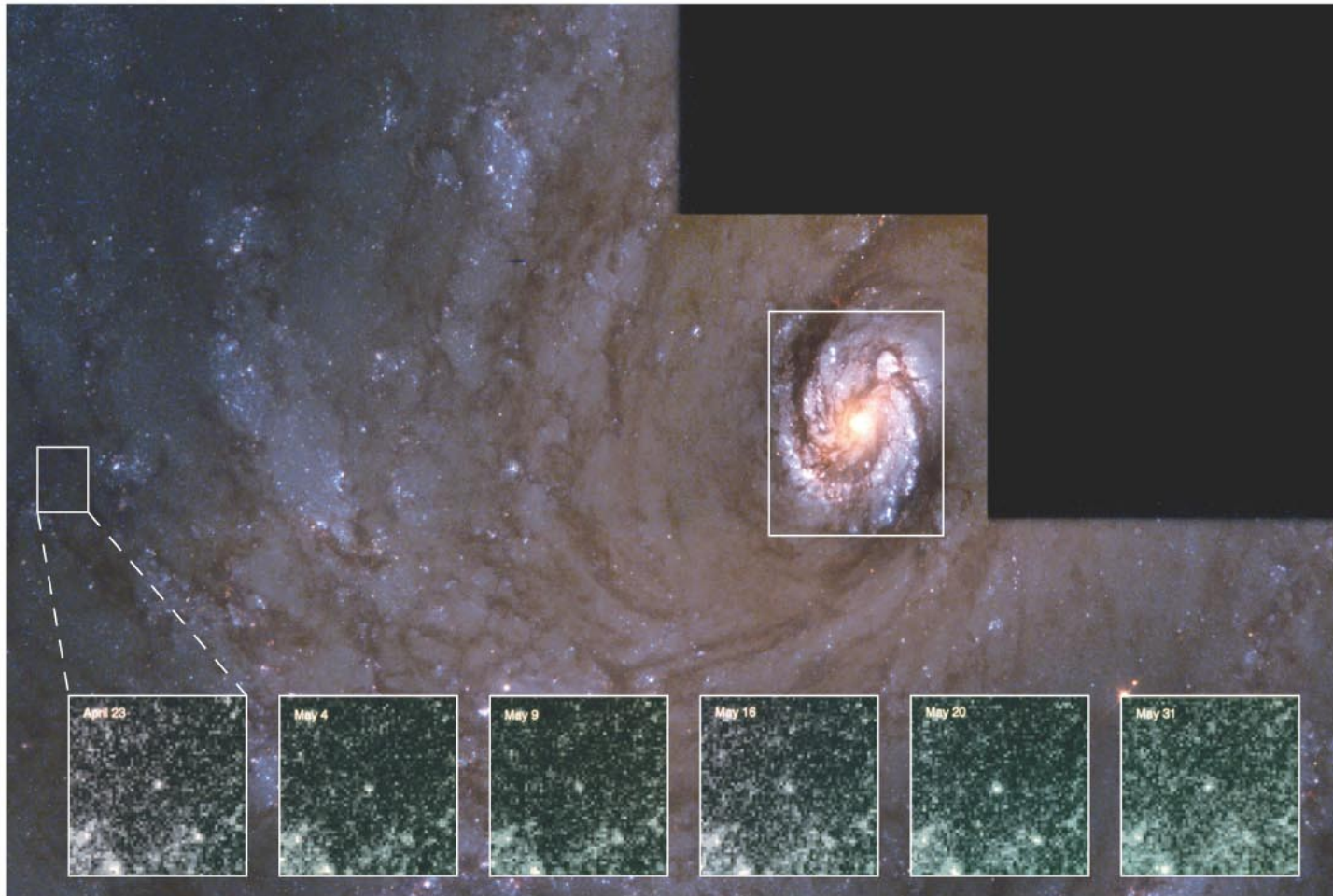
Galaxy Evolution

How did our galaxy (the Milky Way) form? (Ch. 23)



24.2 The Distribution of Galaxies in Space

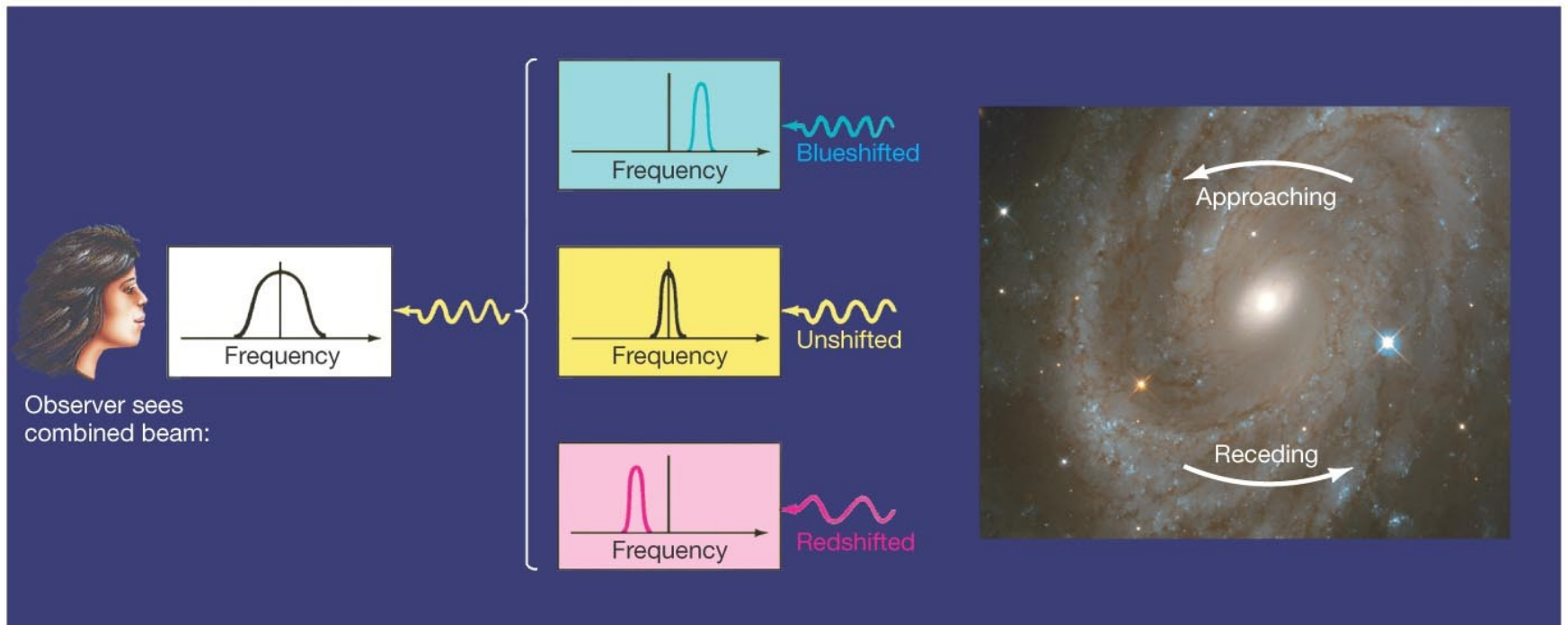
Cepheid variables allow measurement of galaxies to about 25 Mpc away (with HST).



24.2 The Distribution of Galaxies in Space

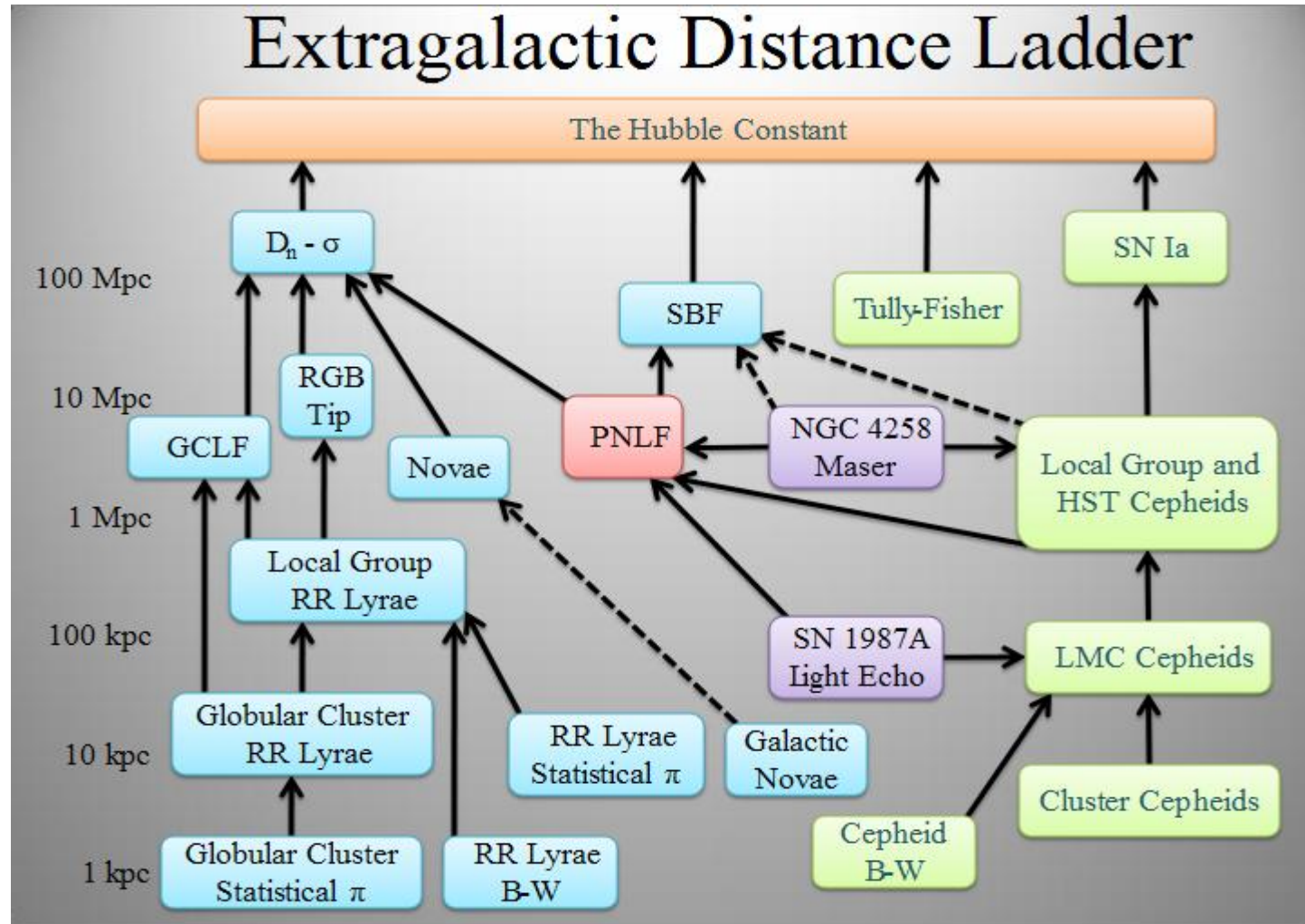
However, some galaxies have no Cepheids, and most are farther away than 25 Mpc. New distance measures are needed.

- **Tully–Fisher relation** correlates a galaxy's rotation speed (which can be measured using the Doppler effect) to its luminosity.



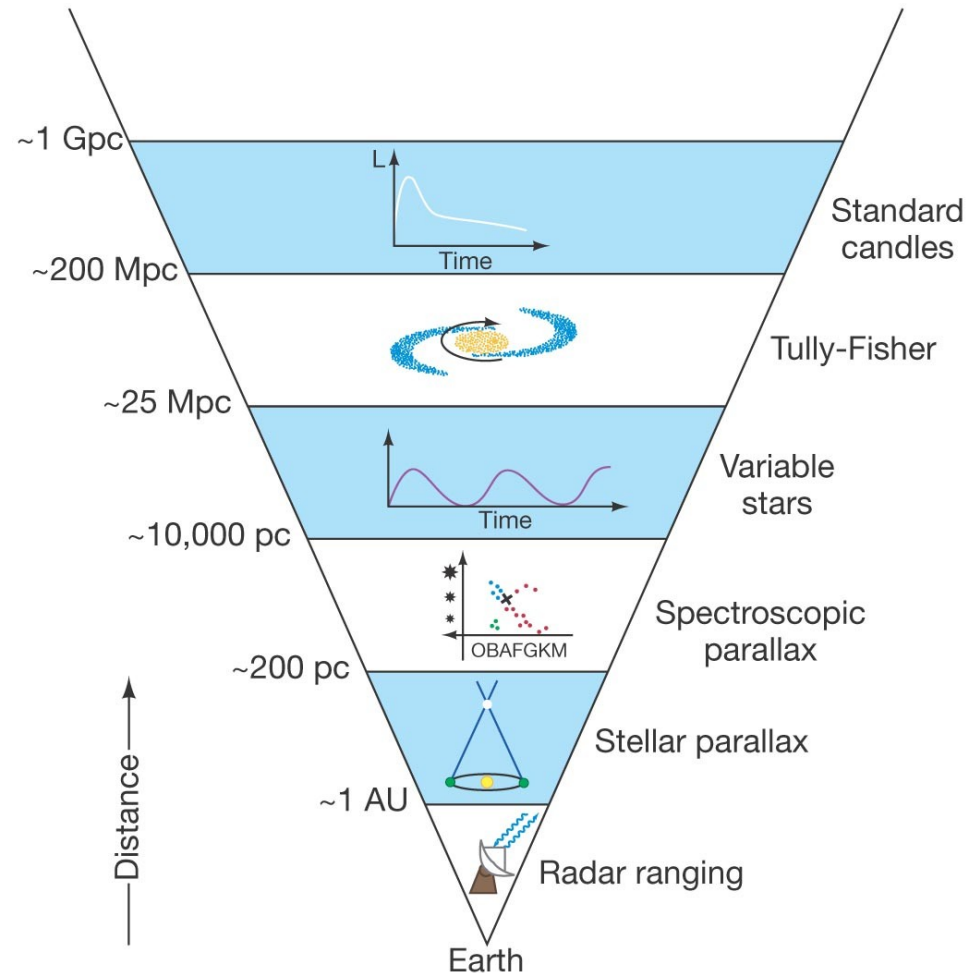
24.2 The Distribution of Galaxies in Space

Actually, about 4 other methods were used to obtain galaxy distances out beyond 30 Mpc. This chart shows how they depended upon other methods.



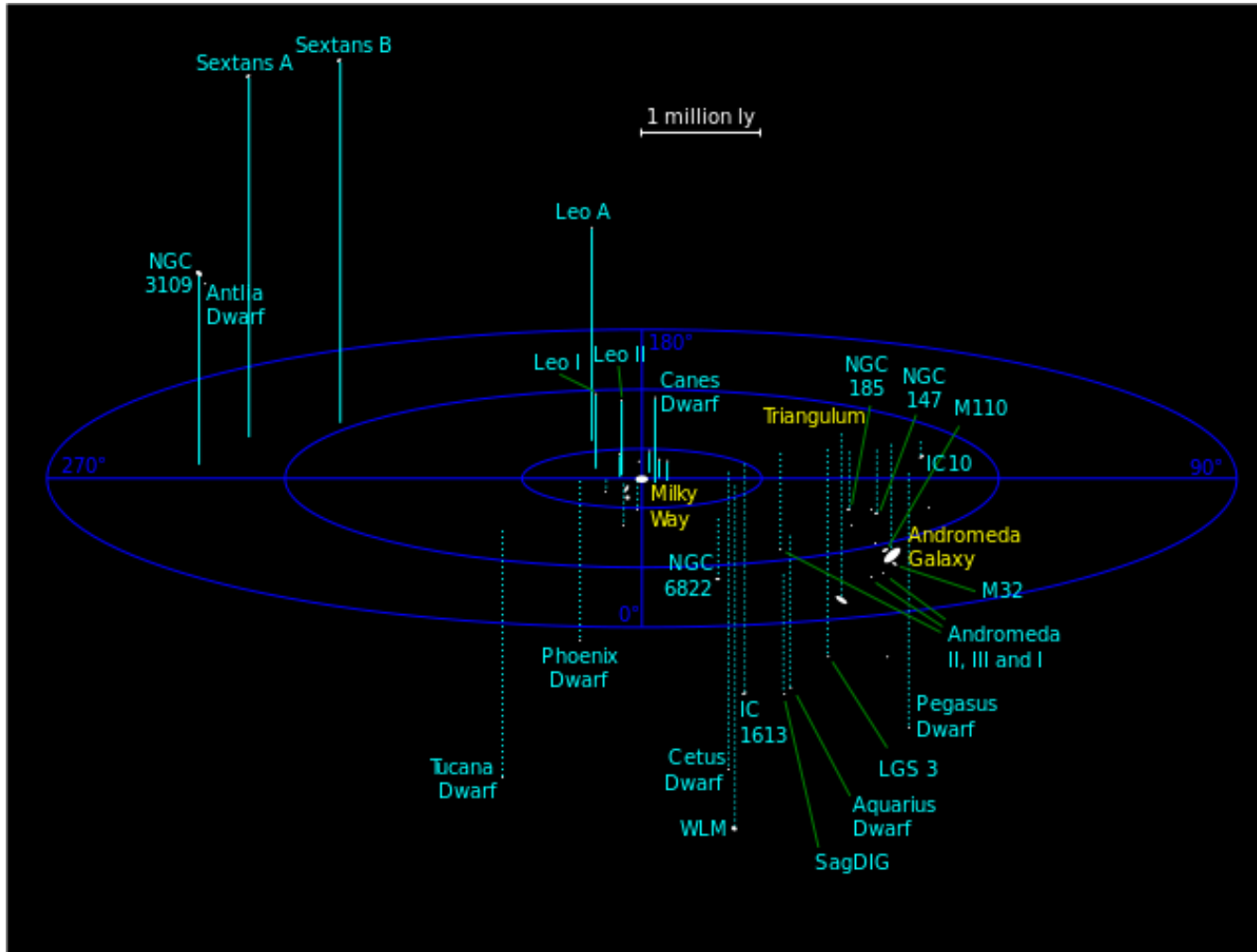
24.2 The Distribution of Galaxies in Space

**With these additions,
the cosmic distance
ladder has been
extended to about 1
Gpc:**



24.2 The Distribution of Galaxies in Space

This is the Local Group of galaxies, about 45 galaxies within about 1 Mpc of the Milky Way.



From wikipedia

24.2 The Distribution of Galaxies in Space

There are three spirals in this group—the Milky Way, Andromeda, and M33. These and their satellites—about 54 galaxies in all—form the Local Group.

Aggregates of galaxies, held together by gravity, are named according to the number of (bright) galaxies roughly as follows:

N=3-10 Poor groups

10-30 Groups

30-50 Poor Clusters

>50 Rich Clusters

24.2 The Distribution of Galaxies in Space

A nearby galaxy cluster is the Virgo cluster; it is much larger than the Local Group, containing over 1500 galaxies.

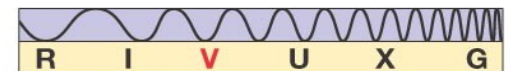
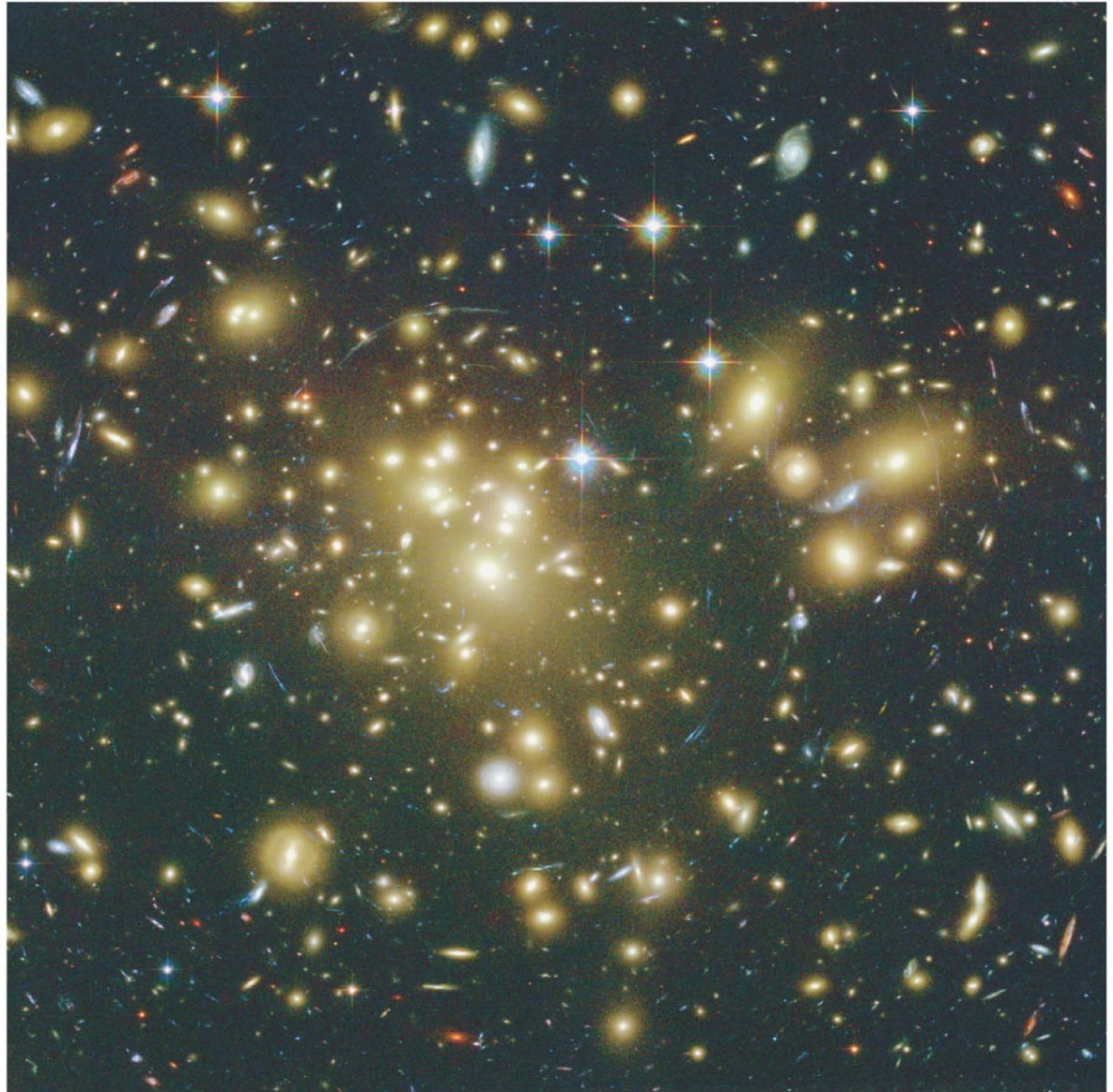
M87



24.2 The Distribution of Galaxies in Space

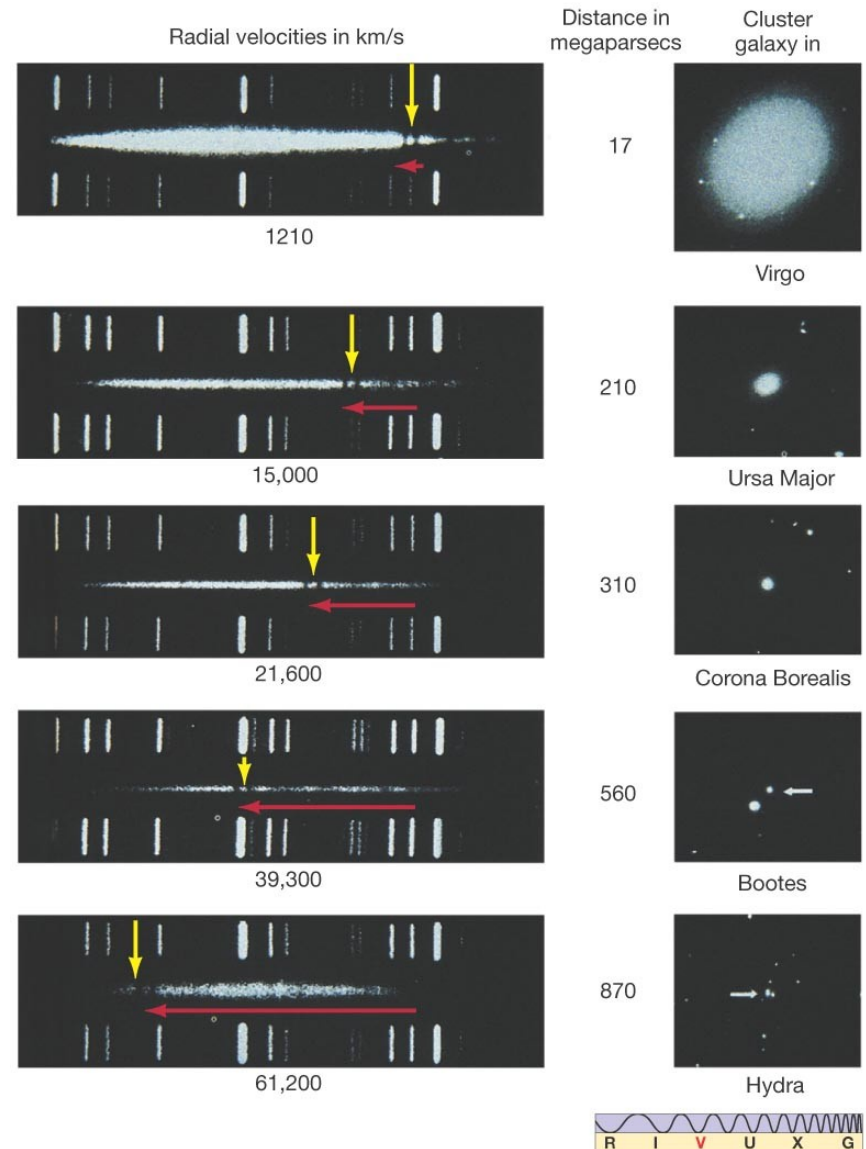
This image shows the Abell 1689 cluster of galaxies, a very large cluster almost 1 billion parsecs away:

See the blue arcs?



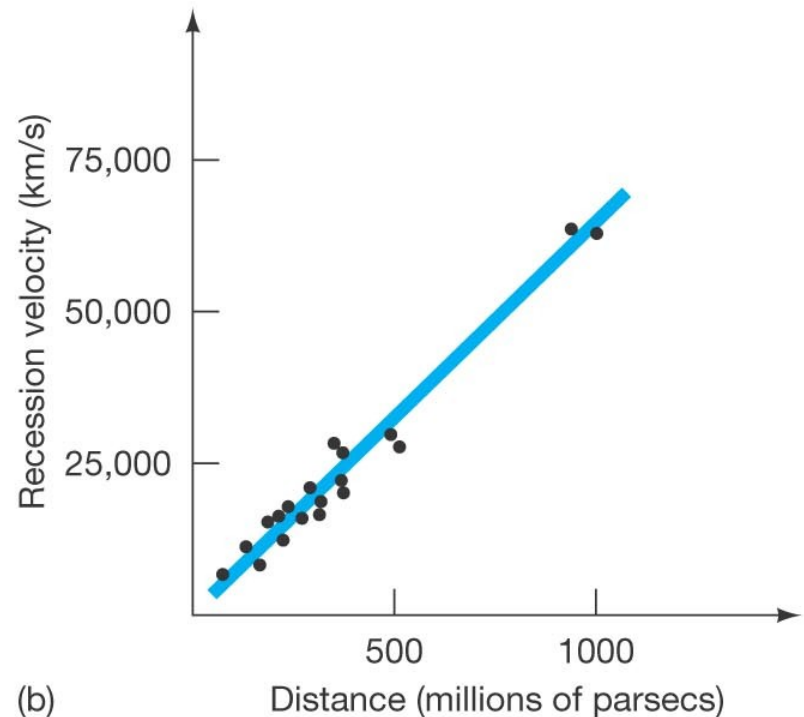
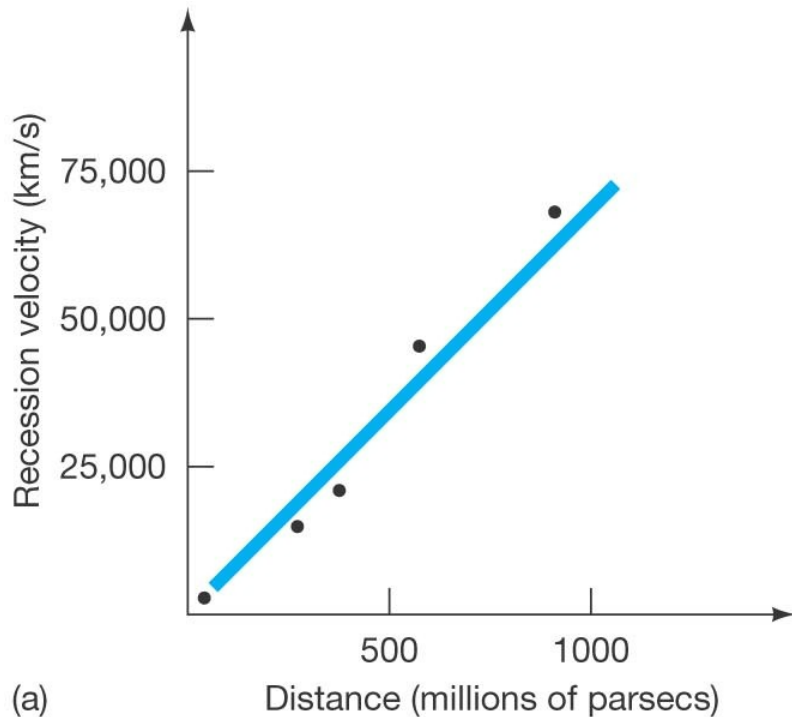
24.3 Hubble's Law

Universal recession: all galaxies (with a couple of nearby exceptions) seem to be moving away from us, with the redshift of their motion correlated with their distance.



24.3 Hubble's Law

These plots show the relation between **distance and recessional velocity** for the five galaxies in the previous figure, and then for a larger sample:



24.3 Hubble's Law

The relationship (slope of the line) is characterized by Hubble's constant H_0 :

$$\text{recessional velocity} = H_0 \times \text{distance}$$

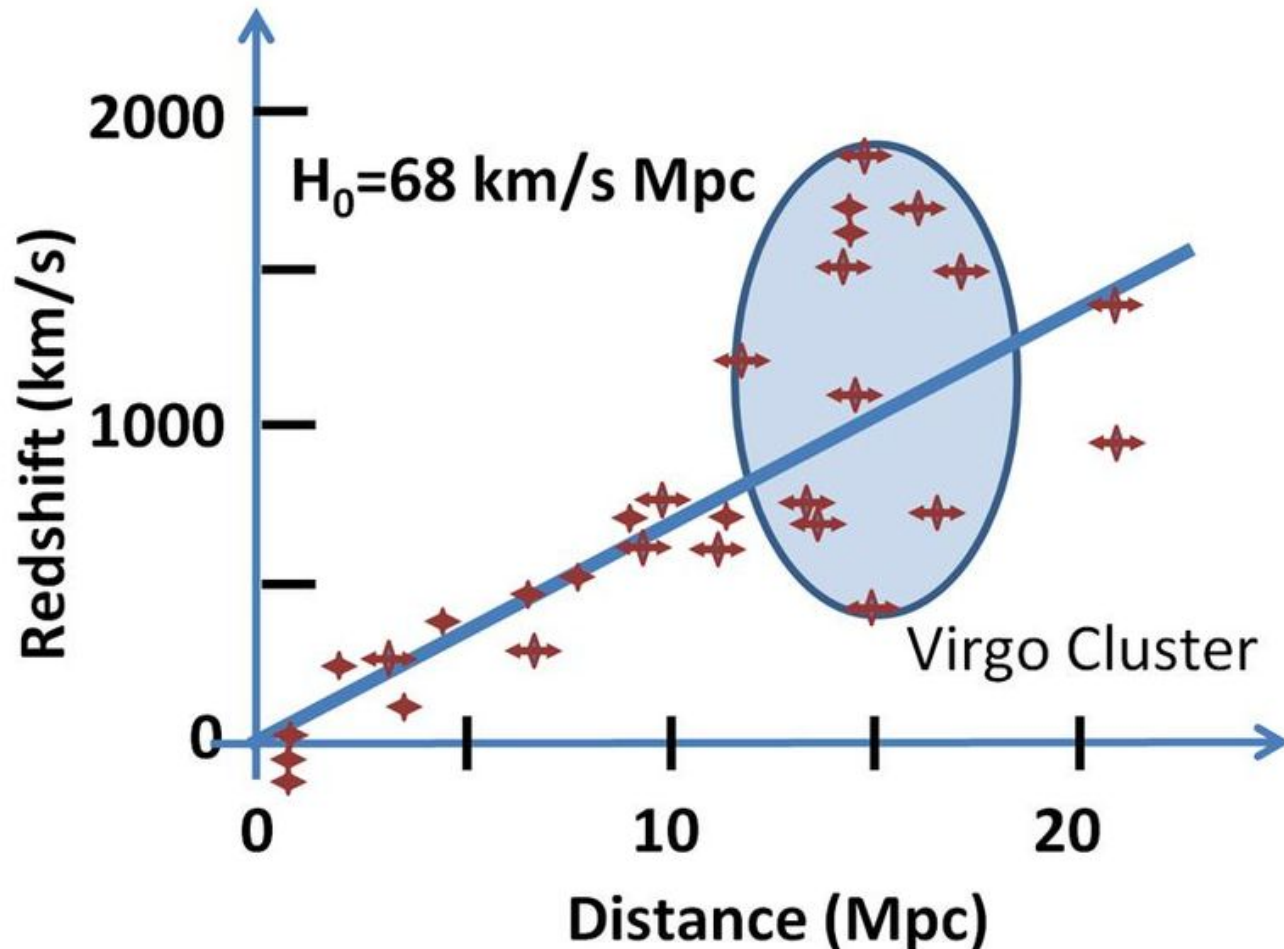
The currently accepted value for Hubble's constant:

$$H_0 = 74 \pm 1.4 \text{ (HST)}, 67.7 \pm 0.4 \text{ km/s/Mpc (Planck)}$$

Measuring distances using Hubble's law actually works better on farther away objects. Nearby, “peculiar motions” overwhelm the expansion velocity.

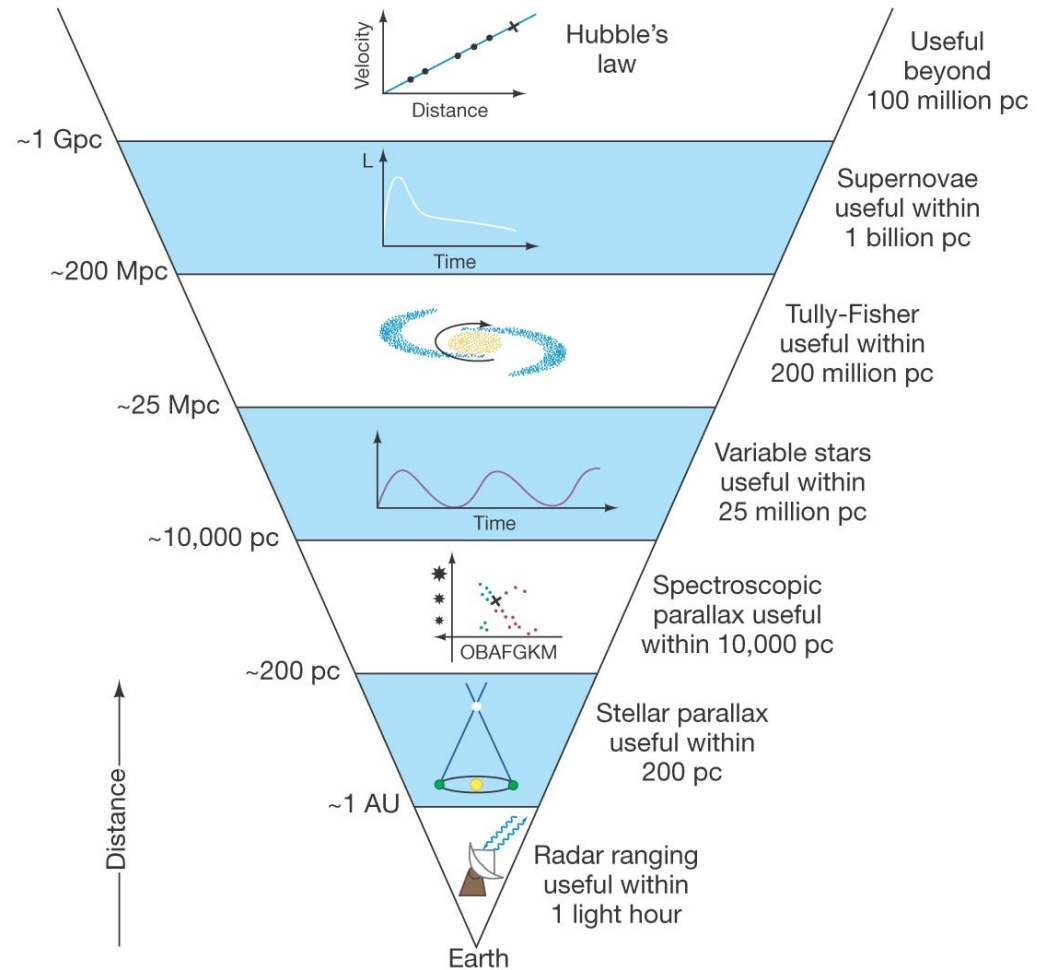
24.3 Hubble's Law

The velocity of a galaxy is also influenced by gravitational attraction to neighboring galaxies!

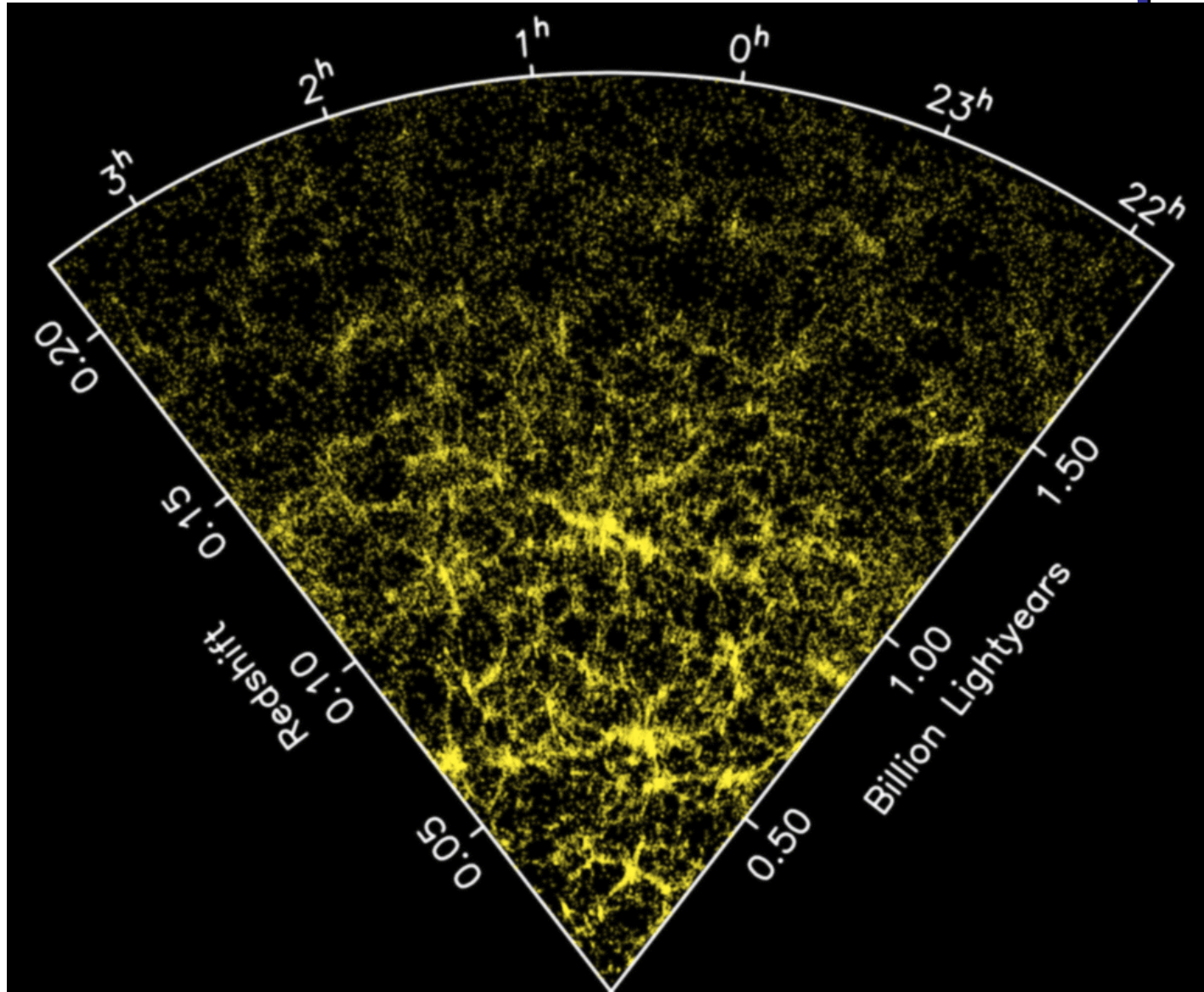


24.3 Hubble's Law

This puts the final step on our distance ladder:



24.2 The Distribution of Galaxies in Space

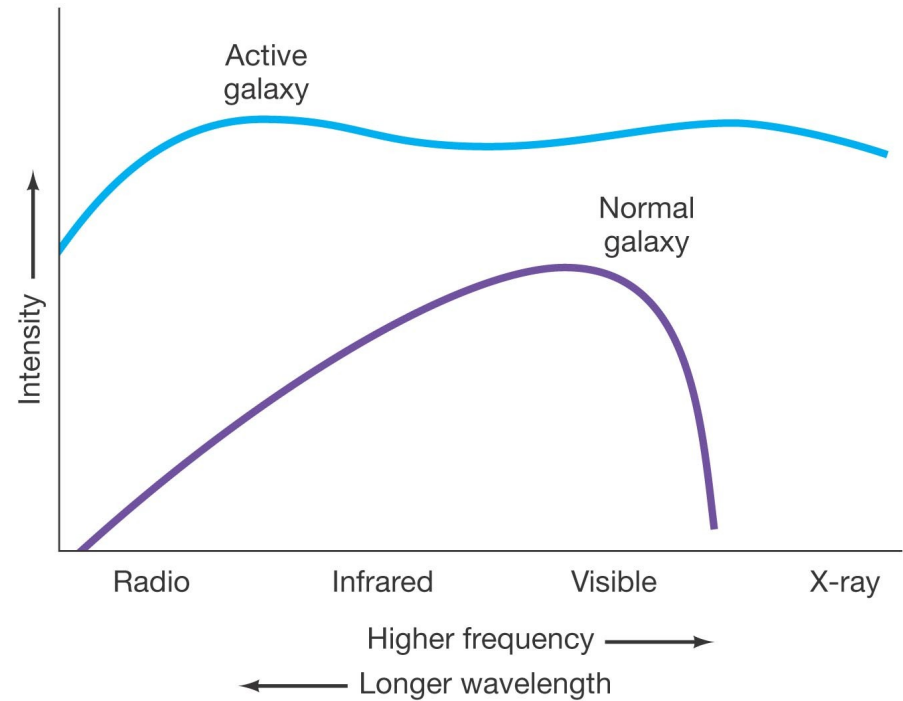


See a “universe in a bottle”

24.4 Active Galactic Nuclei

About 2% of galaxies have unusually **luminous** nuclei.

Such galaxies are called **active galaxies**. They differ from normal galaxies in both the **luminosity** and **type of radiation** they emit:



24.4 Active Galactic Nuclei

Some include starburst galaxies in with the AGN because they have intense star formation going on near their center.

The rest of the AGN are so small that they are unresolved. Types include:

Seyferts

Radio Galaxies (jets and lobes resolvable in the radio)

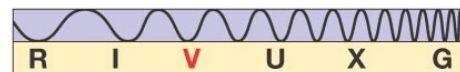
Quasars

Blazars

LLAGN – low luminosity (e.g., LINERS)

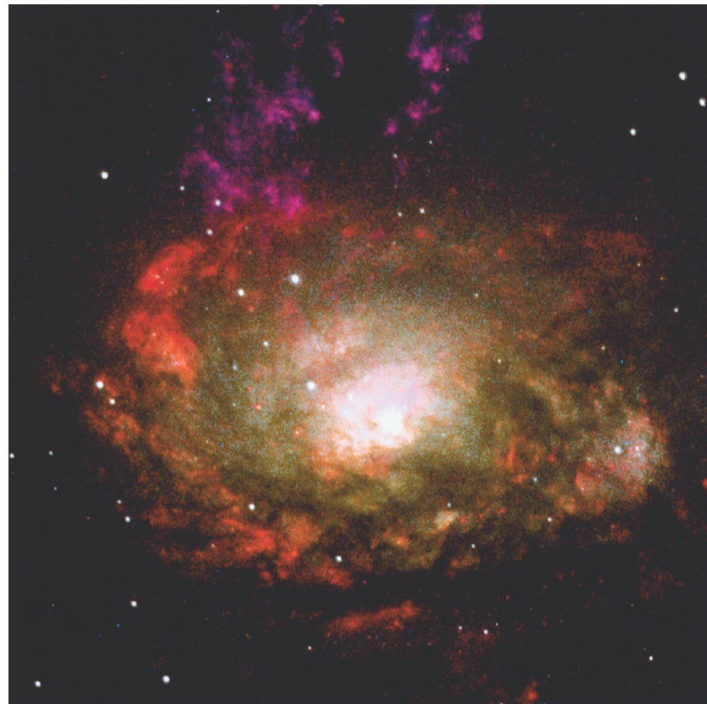
24.4 Active Galactic Nuclei

This active galaxy has star-formation rings surrounding a very luminous core:

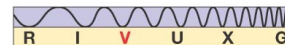


24.4 Active Galactic Nuclei

Seyfert galaxies resemble normal spiral galaxies, but their cores are thousands of times more luminous:

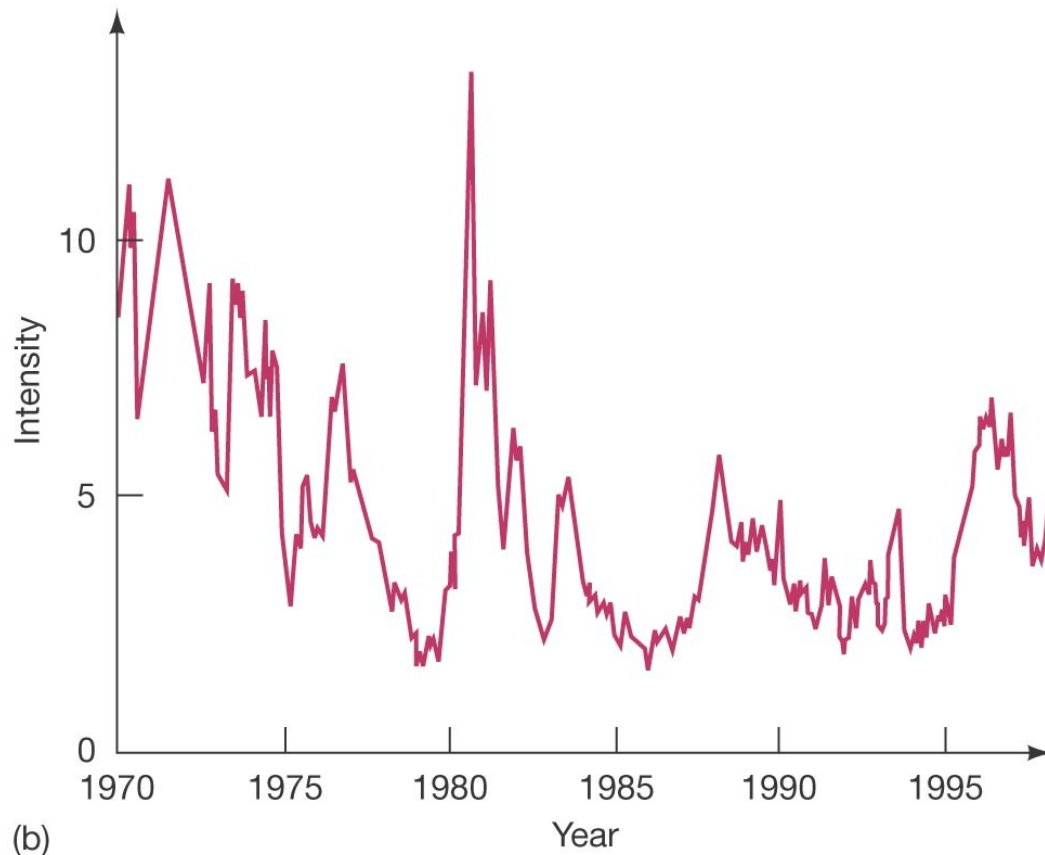


(a)



24.4 Active Galactic Nuclei

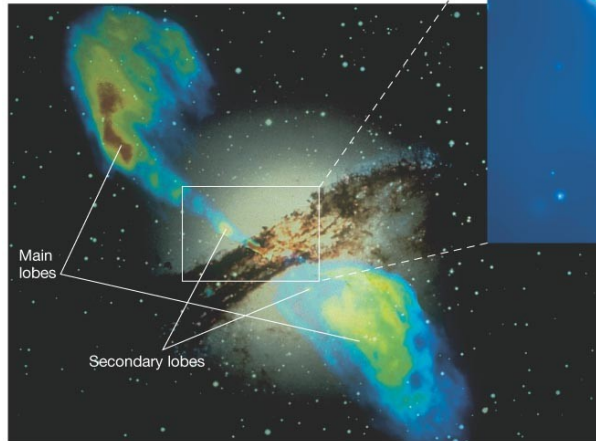
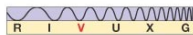
The rapid variations in the luminosity of Seyfert galaxies indicate that the core must be extremely compact:



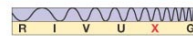
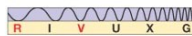
24.4 Active Galactic Nuclei



(a)



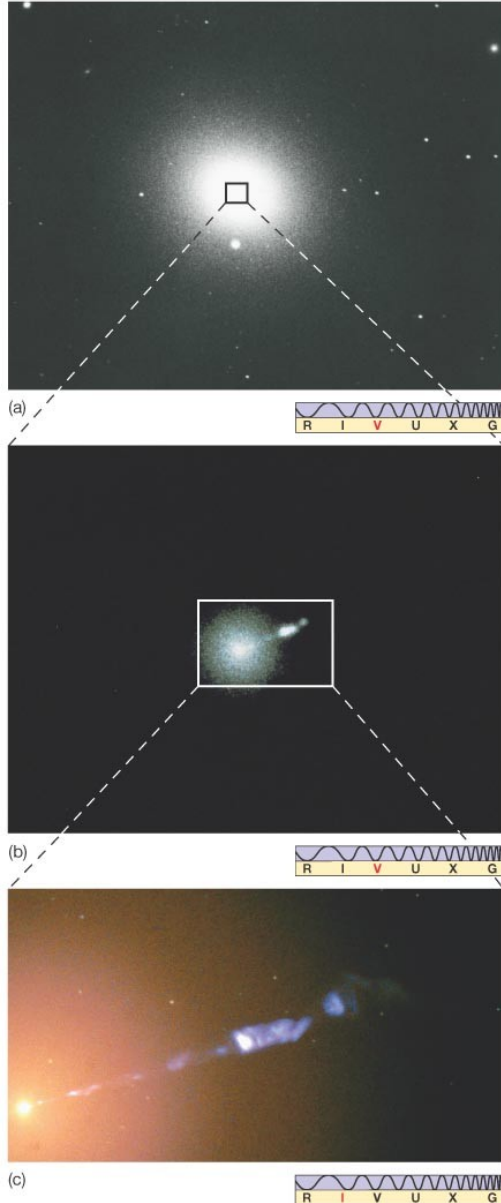
(b)



Radio galaxies emit very strongly in the radio portion of the spectrum. They may have enormous **lobes**, invisible to optical telescopes, perpendicular to the plane of the galaxy.

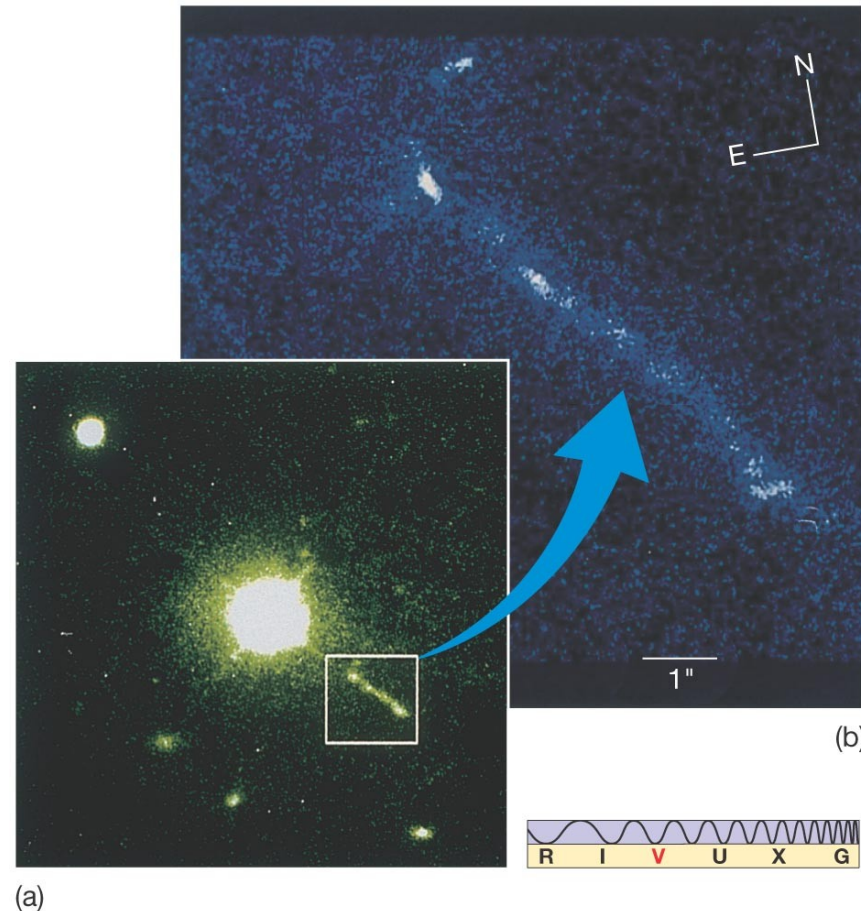
24.4 Active Galactic Nuclei

Many active galaxies have **jets**, and most show signs of interactions with other galaxies.



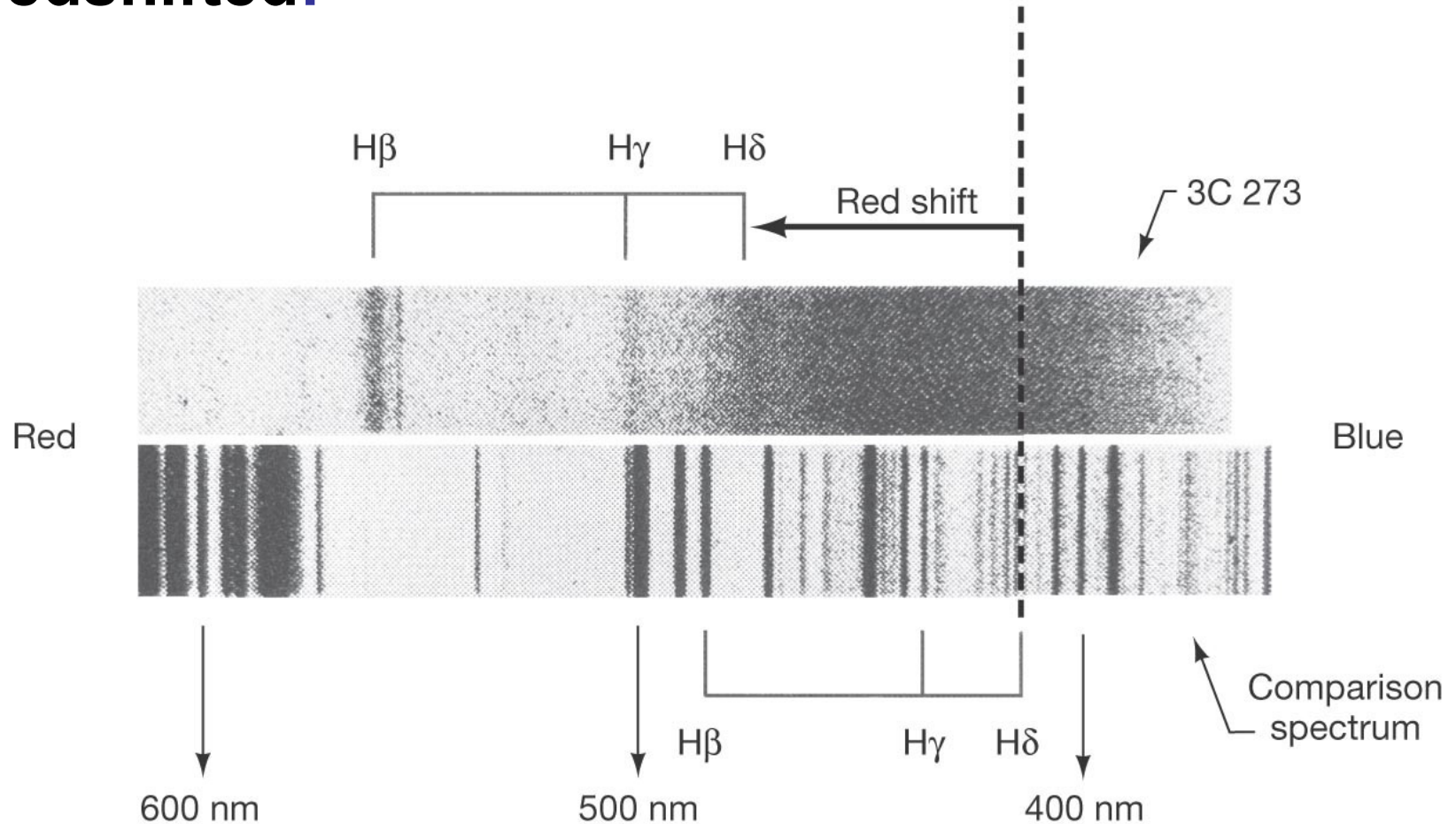
24.4 Active Galactic Nuclei

Quasars—quasi-stellar objects—are starlike in appearance, but have very unusual spectral lines. Some sport optical jets:



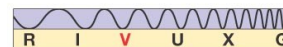
24.4 Active Galactic Nuclei

Eventually it was realized that quasar spectra had known emission lines, but enormously redshifted:

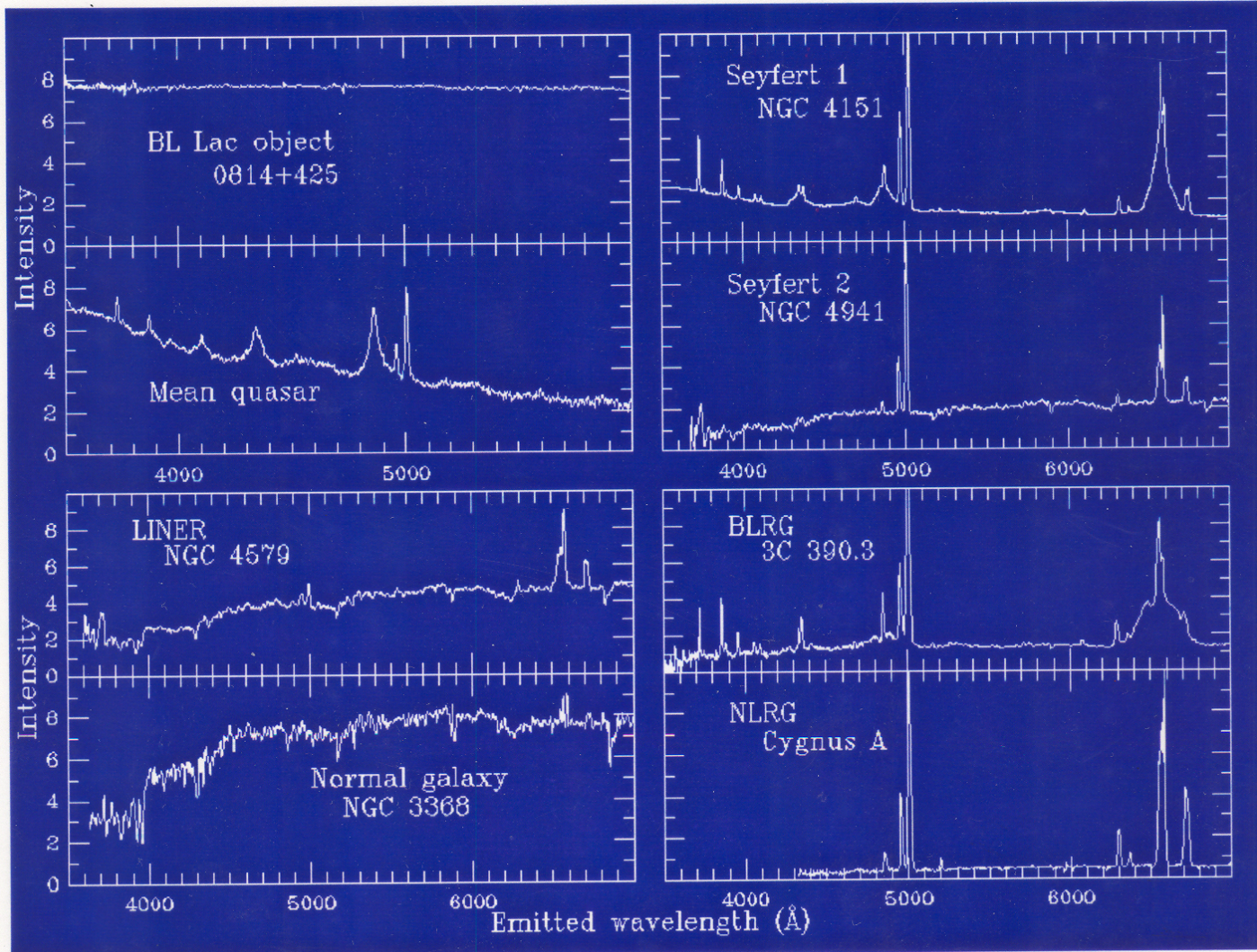


24.4 Active Galactic Nuclei

Solving the spectral problem introduces a new problem—quasars must be among the most **luminous** objects in the Universe, to be visible over such enormous distances.



Optical Spectra of AGN



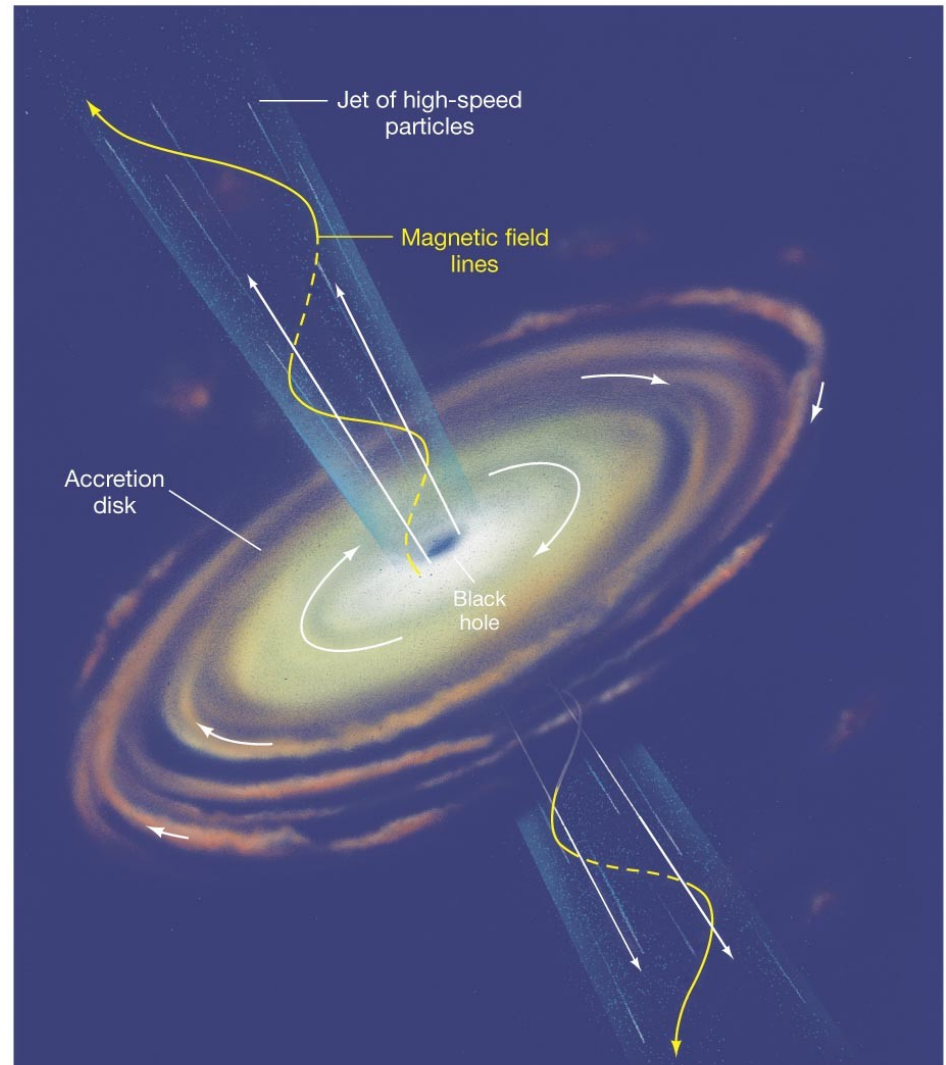
24.5 The Central Engine of an Active Galaxy

Active galactic nuclei have some or all of the following properties:

- **high luminosity**
- **nonstellar energy emission**
- **variable energy output, indicating small nucleus**
- **jets and other signs of explosive activity**
- **broad emission lines, indicating rapid rotation**

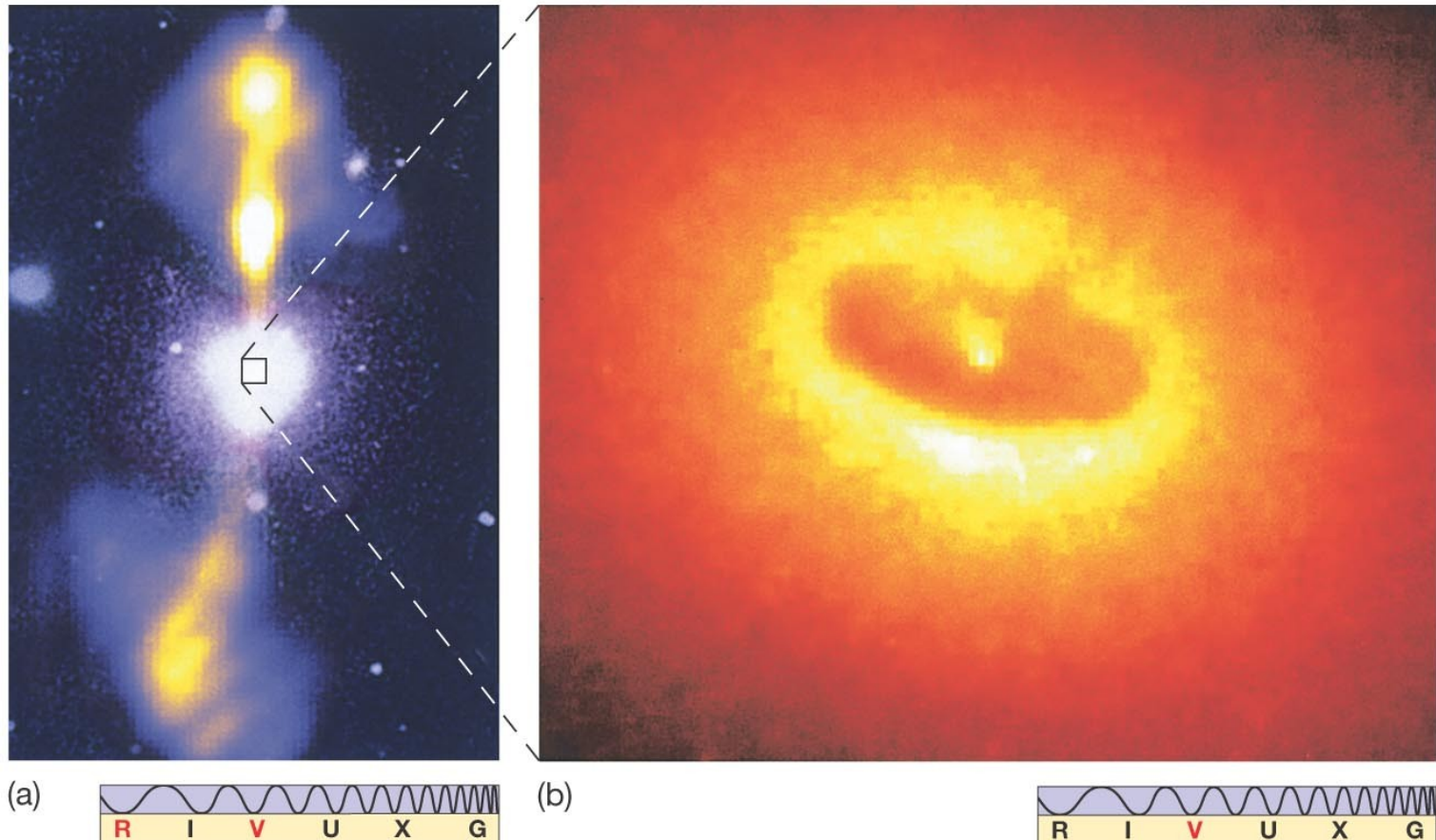
24.5 The Central Engine of an Active Galaxy

This is the leading theory for the energy source in an active galactic nucleus: a **black hole**, surrounded by an **accretion disk**. The **strong magnetic field lines** around the black hole channel particles into **jets** perpendicular to the magnetic axis.



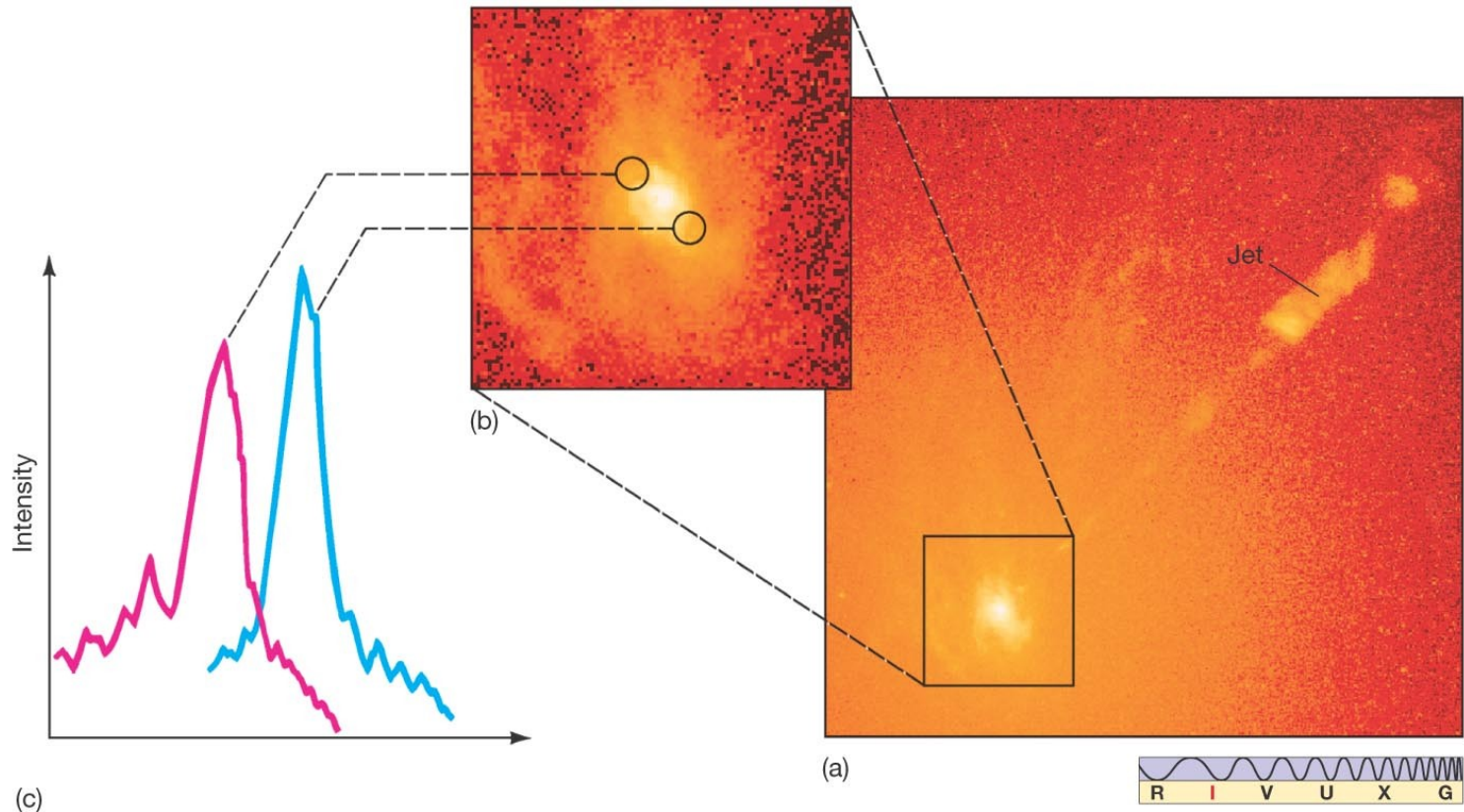
24.5 The Central Engine of an Active Galaxy

This pair of images shows evidence for a black hole at the center of NGC 4261.



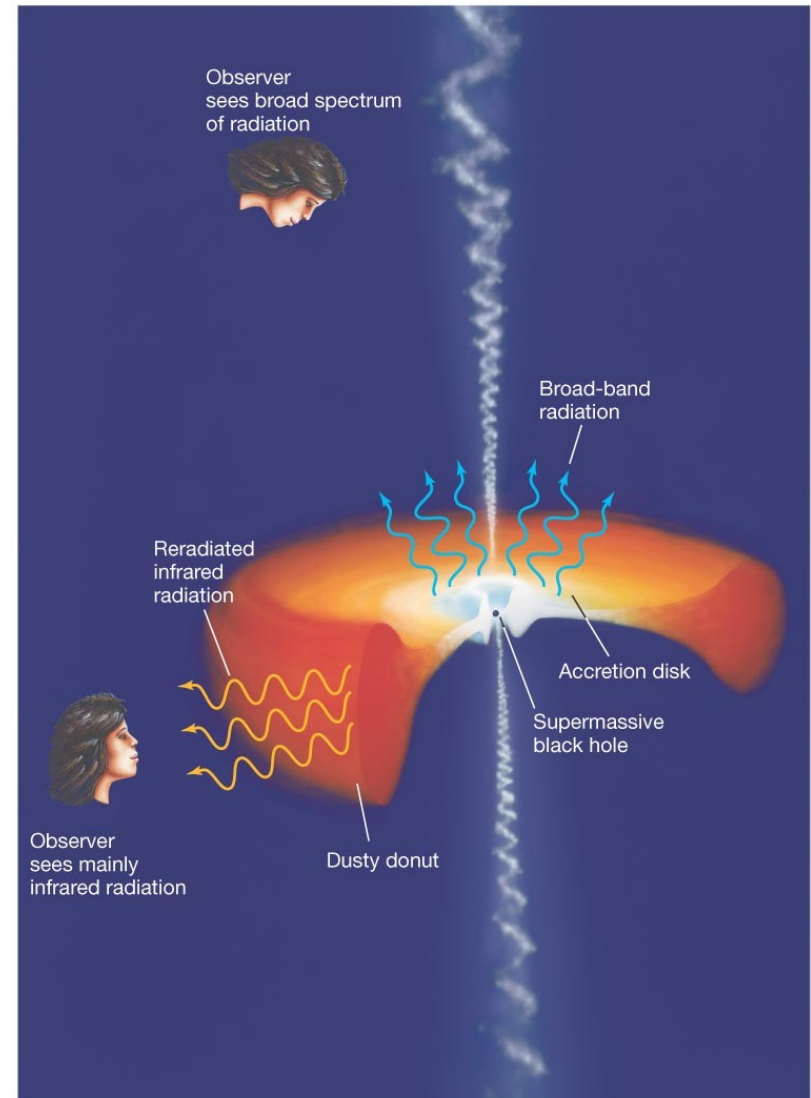
24.5 The Central Engine of an Active Galaxy

The central portion of M87 shows rapid motion and jets characteristic of material surrounding a black hole.



24.5 The Central Engine of an Active Galaxy

One might expect the radiation to be mostly X- and gamma-rays, but apparently it is often “**reprocessed**” in the dense clouds around the black hole and re-emitted at longer wavelengths.



Summary of Chapter 24

- Hubble classification organizes galaxies according to shape.
- Galaxy types: spiral, barred spiral, elliptical, irregular
- Objects of relatively uniform luminosities are called “standard candles”; examples include RR Lyrae stars and Type I supernovae.
- The Milky Way lies within a small cluster of galaxies called the Local Group.
- Other galaxy clusters may contain thousands of galaxies.

Summary of Chapter 24 (cont.)

- **Hubble's Law: Galaxies recede from us faster the farther away they are.**
- Active galaxies are far more luminous than normal galaxies, and their radiation is nonstellar.
- Seyfert galaxies, radio galaxies, and quasars all have very small cores; many emit high-speed jets.
- Active galaxies are thought to contain supermassive black holes in their centers; infalling matter converts to energy, powering the galaxy.